The dependence of the $A_V$ prior for SN Ia on host mass and disc inclination

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ABSTRACT

Type Ia supernovae (SNe Ia) are used as ‘standard candles’ for cosmological distance scales. To fit their light-curve shape–absolute luminosity relation, one needs to assume an intrinsic colour and a likelihood of host galaxy extinction or a convolution of these, a colour distribution prior. The host galaxy extinction prior is typically assumed to be an exponential drop-off for the current supernova programmes ($P(A_V) \propto e^{-A_V/\tau_0}$). We explore the validity of this prior using the distribution of extinction values inferred when two galaxies accidentally overlap (an occulting galaxy pair). We correct the supernova luminosity distances from the SDSS-III supernova projects (SDSS-SN) by matching the host galaxies to one of three templates from occulting galaxy pairs based on the host galaxy mass and the $A_V$-bias–prior-scale ($\tau_0$) relation from Jha et al. We find that introducing an $A_V$ prior that depends on host mass results in lowered luminosity distances for the SDSS-SN on average but it does not reduce the scatter in individual measurements. This points, in our view, to the need for many more occulting galaxy templates to match to SN Ia host galaxies to rule out this possible source of scatter in the SN Ia distance measurements. We match occulting galaxy templates based on both mass and projected radius and we find that one should match by stellar mass first with radius as a secondary consideration. We discuss the caveats of the current approach: the lack of enough radial coverage, the small sample of priors (occulting pairs with HST data), the effect of gravitationally interacting as well as occulting pairs, and whether an exponential distribution is appropriate. Our aim is to convince the reader that a library of occulting galaxy pairs observed with HST will provide sufficient priors to improve (optical) SN Ia measurements to the next required accuracy in cosmology.

Key words: dust, extinction – galaxies: general – galaxies: ISM – galaxies: structure – distance scale.

1 INTRODUCTION

Type Ia supernovae (SNe Ia) are a prime cosmological tool (Riess et al. 1998; Perlmutter et al. 1999). Their relation between light curve and intrinsic luminosity makes them excellent standard candles. Part of the exquisite precision required is a good understanding of the expected dust extinction and known reddening-dimming relation, the extinction law (Cardelli, Clayton & Mathis 1989; Calzetti, Kinney & Storchi-Bergmann 1994). With the dark energy detection result in hand, the supernova community now looks for the next step in accuracy (e.g. the 3 per cent solution; Riess et al. 2011). A much better model for the dust extinction in the supernova’s host galaxy will be critical. Host galaxy dust is already identified by the Dark Energy Task Force as a principal unknown (Albrecht et al. 2006) and this remains the case, together with the photometric uncertainty across surveys (e.g. Conley et al. 2011; Betoule et al. 2014; Scolnic et al. 2014).

These issues are being addressed by the current generation of supernovae searches, e.g. the Supernova Legacy Survey (Astier et al. 2006), the Equation of State: SupErNovae trace Cosmic Expansion (Wood-Vasey et al. 2007), and the Sloan Digital Sky Survey SuperNova survey (SDSS-SN; Kessler et al. 2009; Sako et al. 2014).

A prime candidate is the correlation with properties of the galaxy hosting the SN Ia (Gallagher et al. 2005; Mannucci, Della Valle &
Panagia 2006; Sullivan et al. 2006; Kistler et al. 2013; Pan et al. 2014). Two populations of SNe Ia appear associated with two populations of stars: blue, star-forming galaxies host higher rates of fast declining supernova and red, passive galaxies host predominantly more slowly-declining SNe Ia (Hamuy et al. 1996; Howell 2001; van den Bergh, Li & Filippenko 2005; Mannucci, Della Valle & Panagia 2006; Schawinski 2009; Lampeitl et al. 2010; Wang et al. 2013). The difference in light-curve characteristics may be completely attributable to the different progenitor stellar populations.

However, simultaneously, the issue of host galaxy extinction has come to the fore, either the applicability of the appropriate extinction law (Riess, Press & Kirshner 1996; Phillips et al. 1999; Altavilla et al. 2004; Reindl et al. 2005; Conley et al. 2007; Jha, Riess & Kirshner 2007), the validity of the prior of extinction values (Jha et al. 2007; Wood-Vasey et al. 2007; Gupta et al. 2011), or the possibility that the geometry of the dusty interstellar medium (ISM) in (and hence the extinction distribution) may change with galaxy evolution (Holwerda 2008; Holwerda et al. 2015a). Host galaxies appear to be indistinguishable from normal field galaxies (Childress et al. 2013). After photometric calibration, this host galaxy extinction issue is likely to return as a dominant issue standing in the way of the community’s final goal of 1 per cent precision for SN Ia (e.g. Riess et al. 2011; Kelly et al. 2015).

There is now an opportunity to probe the dependences of the extinction prior through two projects. First, the SDSS-SN project has generated a uniform sample of SNe Ia and their host galaxy properties of exquisite quality. Secondly, occulting galaxy pairs promise to map extinction in spiral discs using the Hubble Space Telescope (HST) imaging, enabling a way to generate an extinction prior for SN Ia fit. The first of these extinction maps are now available.

In this paper, we explore the effects of applying the extinction priors from occulting galaxy pairs to the SDSS-SN sample and gauge the effects. We opt for the SDSS-SN sample among all the supernova samples available now as it has uniform properties of not only the supernovae, as measured with both light-curve packages, but also on the host galaxies. The latter is important if one wants to correct for the effects of host galaxy properties. The SDSS-SN project uses two light-curve fitting packages, MLCS2k2 and SALT2. We use the former’s output as it uses an explicit prior for host galaxy extinction, making a direct comparison with the occulting galaxy technique practical, even though SALT and its derivative BASALT are now often the preferred package (e.g. Guy et al. 2010; Be-toule et al. 2014; Rest et al. 2014). We should note here that the $A_V$ value used in MLCS2k2 is not the exact same quantity as measured from galaxy transparency measurements, i.e. occulting galaxy measurements. MLCS2k2’s $A_V$ parameter is the difference in supernova colour from an assumed intrinsic one, translated to an extinction via an assumed reddening law (in the case of the SDSS-SN; Kessler et al. 2009, a greyer one than Milky Way). The reddening law in this instant is an implicit input setting, one which has been tweaked to minimize cosmological fit residuals (e.g. Hicken et al. 2009).

We assume in this paper that the shape of the $A_V$ fit parameter’s prior and the shape of the distribution of $A_V$ values observed in occulting galaxies is to first order the same.

We explore the extinction effects by matching SN Ia host galaxies to those occulting galaxies with HST data through their stellar mass. In a previous paper on SN Ia (Holwerda et al. 2015a), we explored the relation between the extinction distribution and inclination and found that an inclination correction using a simple $\cos(i)$ is sufficient to make the $A_V$ distribution found by the SDSS-SN identical in high and low-inclination subsamples. We therefore adopt this inclination correction in this paper as well. We will discuss the caveats and limits of the current approach and sketch out how to proceed from here in the generation of improved extinction priors, tailored to the host galaxy of the observed supernova.

\section{OCCULTING GALAXIES}

One can accurately measure the extinction of light by interstellar dust using a known light source. The most accurate extragalactic method is to use a partially occulting galaxy pair (Fig. 1). Assuming both galaxies are symmetric, one can estimate flux contributions to the overlap region from the complementary regions at the same radius in each galaxy. The reliance on differential photometry enables an extinction measurement ($A_V$) in a single-colour image without relying on any assumed extinction law. With multiple filters, the effective extinction law itself can be measured (Keel & White 2001a; Holwerda et al. 2009; Keel et al. 2014).

\subsection{The occulting galaxies method}

Fig. 1 shows an example pair, serendipitously imaged with HST for the ANGST survey (Dalcanton et al. 2009), near NGC 253 (2MASX). To determine the extinction in the overlap region (black aperture), one estimates the relative flux contributions by both galaxies from the complementary apertures in the foreground (green aperture) and background galaxy (red aperture in Fig. 1), respectively, assuming rotational symmetry (i.e. the image can be rotated and be self-similar). Thus, we estimate the flux contributions to the overlap region from the same image. In the case of ground-based images, each of these apertures would yield single $A_V$ value, averaged over a large portion of the foreground disc. In the case of an HST image, it results in a highly accurate map of the dust extinction in the overlap region with hundreds of independent lines of sight. The map is constructed pixel by pixel; for each pixel in the overlap aperture, the corresponding pixel in the background galaxy aperture (red) is found, i.e. the pixel at the same distance from the background galaxy centre but in the opposite direction, as well as the corresponding pixel in the foreground galaxy (same distance with respect to the foreground galaxy centre). We estimate the optical depth ($\tau$) or extinction ($A_V$) from the flux in the overlap ($F + Be^{-\tau}$), a mix of flux from background and foreground galaxies, the flux in the corresponding background pixel ($\bar{B}$) and foreground pixel ($\bar{F}$, same colours as Fig. 1):

\begin{equation}
A_V = 1.086 \times \bar{\tau} = -1.086 \times \ln \left( \frac{(F + Be^{-\tau}) - \bar{F}}{\bar{B}} \right). \tag{1}
\end{equation}
The properties of the foreground occulting galaxy in three pairs with sufficient HST information to form an extinction map. Stellar mass in solar masses and radii in effective (half-light) radii.

<table>
<thead>
<tr>
<th>Pair</th>
<th>Type</th>
<th>Stellar mass</th>
<th>Radial coverage</th>
</tr>
</thead>
<tbody>
<tr>
<td>UGC 3995</td>
<td>Sa</td>
<td>$1.7 \times 10^{10}$</td>
<td>0–2</td>
</tr>
<tr>
<td>2MASX</td>
<td>Sd</td>
<td>$3.1 \times 10^{9}$</td>
<td>2–5</td>
</tr>
<tr>
<td>AM 0500</td>
<td>Sbc</td>
<td>$2.6 \times 10^{9}$</td>
<td>$\sim 3.7$</td>
</tr>
</tbody>
</table>

and that the colour–extinction relation is grey in the ground-based observations (Holwerda et al. 2007). The Galactic extinction law returns as soon as the physical sampling of the overlap region resolves the molecular clouds in the foreground disc (<100 pc; Elmegreen et al. 2001; Keel & White 2001a,b; Holwerda et al. 2009).

The extinction maps of foreground discs made with HST are especially useful as a probability map of extinction towards a single object of interest. To date, we have obtained HST imaging of three overlapping pairs with a large enough overlap region to generate a reasonable distribution. The normalized distributions of these three are shown in Fig. 2. Radial coverage is very similar to the radii at which SNe Ia are found (Table 1): between 0 and 6 effective radii (converted using the prescriptions in Graham & Driver 2005).

Our previous results on occulting galaxy pairs are consistent with an exponential decline with radius of the average disc attenuation, not dissimilar to the exponential profile of the stellar light (e.g. Domingue et al. 1999; Holwerda 2005; Holwerda et al. 2005a,b). We stress here that we deal here now in the exponential distribution of values in a certain range of radii but this is not the same as the radial exponential decline. For example, the lower values ($A_V \sim 0.1$) are typically, but not exclusively, found at higher radius while the higher values are more typical for spiral arms.

2.3 UGC 3995

This pair of galaxies has been long known to be overlapping and interacting (e.g. Marziani et al. 1999) and we have analysed the extinction properties of both CALIFA DR1 IFU data and archival WFPC2 F606W imaging in Holwerda & Keel (2013). Radial coverage is from 5–15 arcsec, covering the full 15.60 arcsec Petrosian radius (SDSS-r) of UGC3995B, the foreground galaxy. We obtained the foreground galaxy’s mass from the SDSS fluxes of the unoccupied side and the prescription from Zibetti, Charlot & Rix (2009).

2.4 2MASX

This pair’s serendipitous discovery was originally reported in Holwerda et al. (2009). The distance and stellar mass was estimated in Holwerda et al. (2013b) and the radial coverage of the overlap region is between 2 and 5 effective (or half-light) radii ($R_e$).

2.5 AM 500–620

This was a pair identified in White et al. (2000) as an overlapping galaxy pair and the HST data were originally reported in Keel & White (2001a). The radial coverage is mostly in the outer disc ($\sim R_{25}$, the 25 mag arcsec$^{-2}$ as defined by de Vaucouleurs et al. 1991). Stellar mass and Hubble type are from Keel & White (2001a).
3 TYPE IA SUPERNOVAE SAMPLE

SNe Ia are the touchstone of cosmological distance measurements. The photometric properties of their light curves are very stable but appear to depend slightly on host galaxy properties (e.g. star formation rate or stellar mass). With this mind, the third incarnation of the SDSS (SDSS-III) included a supernova search with spectroscopic follow-up. This SDSS-Supernova search (SDSS-SN)\(^1\) has yielded a wealth of galaxy and supernova properties.

SDSS-SN is described in Frieman et al. (2008) and in more detail in Sako et al. (2008) and Sako et al. (2014). Observing for three months of the year from 2005 to 2007, SDSS-SN identified hundreds spectroscopically confirmed SNe Ia in the redshift range 0.05 < \(z\) < 0.35.

Sako et al. (2014) present the final data-product of this tremendous observational effort. They include an assessment of SN type based on the light curve (PSNID output) and light-curve fits using the two most commonly used packages, SALT2 (Guy et al. 2007) and MLCS2k2 (Jha et al. 2007). Of these two packages, the SALT2 light-curve fitter is now the most popular. Yet we focus on the MLCS2k2 package because it explicitly starts with a host galaxy extinction prior, making it easier to directly compare its assumptions to our observations in occulting pairs.

However, we should note that both SALT2 and MLCS2k2 essentially measure the reddening of the supernova, not actual attenuation but the convolution of intrinsic supernova colour distribution, scatter by the surrounding interstellar matter and dust attenuation in the host galaxy. The shape of the MLCS2k2 \(A_V\) prior used for the light-curve fits was \(P(A_V) \propto e^{-A_V/\tau_0}\), with \(\tau_0 = 0.4\). Because it has host galaxy extinction as an explicit prior and a model relation between \(A_V\) distribution width (\(\tau_0\)) and \(A_V\) bias, we use the MLCS2k2 values for our further analysis here. The MLCS2k2 version used for the SDSS-SN analysis assumed an \(R_V = 2.8\) reddening relation to convert reddening to extinction \(A_V\) (Kessler et al. 2009), indicating it is not completely host attenuation alone. However, it is not as unphysical as the \(R_V = 1.7\) needed elsewhere with MLCS2k2 (Hicken et al. 2009). For our purposes here, we will explicitly assume that the shape of the prior for MLCS2k2’s fit parameter \(A_V\) can be taken to be similar to the distribution of actual attenuation values as found in occulting galaxies.

Galaxy properties are those available from the SDSS-DR9 data base and stellar mass and star formation are modelled with two different packages, FSFS (Conroy & Gunn 2010) and PEGASE (Bruzual A. 2009). For the purpose of this paper, the stellar mass is of interest and we choose the FSFS value for the further analysis. Fig. 3 shows the distribution of host mass galaxies with SN Ia. The mean mass is \(\sim 10^{10}\,\text{M}_\odot\), similar to AM500. We focus only on those objects that have reasonable chance of being bona fide SNe Ia (S/N > 5 and \(P(\text{SN Ia}) > 50\) per cent) to ensure the conclusions for the improved prior are based on these only (1698 supernovae in the total sample).

In addition to the values thoughtfully provided by the SDSS-SN project in Sako et al. (2014), we retrieve the axis ratio, position angle, and Petrosian radius from the SDSS server. Using these values, we compute the galactocentric radial distance from the centre of the host galaxy in Petrosian radii. This is a different value than that of the separation presented in Sako et al. (2014), the ‘directional light radius’ \((d_{\text{MLR}})\), which is not deprojected into the plane of the host disc. If we assume discs \((n=1\,\text{Sérsic index})\), the Petrosian radius corresponds to approximately two effective radii \((R_e = 2.15 \times R_e)\).

Figure 3. The distribution of SN Ia host galaxy mass from Sako et al. (2014). They present both FSFS (Conroy & Gunn 2010) and PEGASE (Bruzual A. 2009) stellar masses. The masses of 2MASX (thick dashed), UGC 3995 (dashed line), and AM500 (dotted line) are marked as well.

Figure 4. The distribution of the projected separation of SNe Ia from their host galaxy in effective radii \((R_e)\) based on the values reported in Sako et al. (2014). For comparison, Sako et al. (2014) plot the distance-to-host or ‘directional light radius’ \((d_{\text{MLR}})\), which is not deprojected into the plane of the host disc.

see Graham & Driver 2005). Fig. 4 shows the radial coverage for the SNe Ia for which SDSS Petrosian radii were available. SNe Ia are typically found by SDSS-SN between 0 and 5 effective radii (<2.5 Petrosian radii), which correspond approximately to the radii covered by the occulting galaxy pairs (Table 1).

4 ANALYSIS

4.1 \(A_V\) distribution and bias

The relation between the exponential drop-off and bias introduced in the \(A_V\) value was explored by Riello & Patat (2005) simulations. Fig. 5 shows their relation (fig. 22 in the appendix of Jha et al. 2007). We mark the \(\tau_0\) values inferred from

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\(^1\) http://www.sdss.org/supernova/aboutsupernova.html
Figure 5. The relation between prior exponential scale and the resulting bias in the \( A_V \) from Jha et al. (2007). The thick grey line is the MLCS2k2 prior. The values of the three occulting pairs are also marked (not corrected for their inclination).

the occulting galaxy pairs, corrected for inclination (\( \tau_0 = \tau / \cos(i) \)). To parametrize the \( A_V \) bias, we fit the relation in Fig. 5 with

\[
\Delta A = -0.5 \times e^{-\tau_0/0.15} + 0.02 \times \tau_0,
\]

where \( \tau_0 \) is the exponential drop-off of the host galaxy prior under consideration.

We note first that none of our three occulting galaxy templates match the drop-off typically used in MLCS2k2 (\( \tau_0 = 0.4 \)) which was used for the SDSS-SN analysis (Sako et al. 2008, 2014). Therefore, assuming the SDSS-SN galaxies adhere to one of these three templates based on their host mass will inevitably introduce a bias in the distribution. Our goal is to explore in which direction and how severe it would be.

4.2 New luminosity distances for the SDSS-SN

Fig. 6 shows the relation between the luminosity distance and redshift for the SDSS-SN SNe Ia. They closely follow the expected

\[
\Delta D_L = \frac{(b/a)^2 - (b/a)_{\min}^2}{1 - (b/a)_{\min}^2},
\]

originally from Hubble (1926), where \( a \) is the major axis, \( b \) the minor axis and we assume \( b/a_{\min} = 0.1 \). We found in Holwerda et al. (2015b) that the distribution of SDSS-SN \( A_V \) values becomes the same for high- and low-inclination galaxies, once the values are corrected for disc inclination using this simple \( \cos(i) \) correction (i.e. the screen approximation).

Fig. 7 shows the distribution of resulting \( \tau_{\text{host}} \) for the SDSS-SN SNe Ia sample. The peak around \( \tau_{\text{host}} = 0.4 \) corresponds to all the host galaxies for which no mass estimate was available in the SDSS-SN catalogue. Starting from the new \( \tau_{\text{host}} \) exponential values for each host galaxy (equation 3), we can now calculate the \( A_V \) bias for each host, based on its stellar mass and inclination (equation 2) and apply this to the luminosity distance. Fig. 8 shows the relative difference with \( \Lambda \)CDM distances as a function of redshift and the mean values (and dispersion) before our \( A_V \) bias correction and after (Table 2).

There are two things to note in the corrected values. First is that the mean luminosity distance is slightly lower, i.e. there is a little more extinction in the SN Ia host galaxies if we use individual priors, but there is no less scatter in the values. The change in the mean is interesting for cosmological study but at present only points to the direction one can expect the luminosity distances to change. The priors may be slightly more tailored but
they are hardly appropriate for each host galaxy (see below for the occulting galaxy pair caveat). This lack of an appropriate prior is reflected – in our opinion – in the lack of change in dispersion. If an appropriate correction to the prior shape and hence $A_V$ bias had been applied, luminosity distances would adhere closer to the Hubble flow and any remaining scatter can be attributed to individual motion of the galaxies or photometric errors.

Alternatively, one could match the few occulting galaxy templates by their projected radius, not the foreground galaxy’s mass. These are the blue averages points in Fig. 8. The scatter around the mean is worse than in the case of the mass-matched occulting galaxy templates matched with host mass or radius. The central assumption in this work has been that the attenuation seen in occulting galaxies is the same as the $A_V$ parameter in MLCS2k2. However, because the $A_V$ is derived assuming an attenuation law, it cannot be ruled out that some measure of the reddening is not from dust attenuation but scatter in, or supernova intrinsic colour scatter.

5.2 Radial coverage

While Table 1 and Fig. 4 seem to suggest that at least the SN Ia occur mostly at the same radii for which we have some extinction distribution from occulting pairs, the coverage is far from complete. For example, for low-mass galaxies, there is only information between 1–2 effective radii.

5.3 Small sample

It is obvious to point out that it is impossible to infer a relation between stellar mass and an extinction distribution with only three galaxy templates. We plan to obtain more templates from future HST observations (Holwerda 2014) but for now we can only assign the exponential prior to the closest stellar mass to the host mass. Alternatively, we fitted a linear relation between $r_0$ and mass for the occulting galaxy pairs but this resulted in no improvement of the variance of the SN Ia distances.

5.4 Interactions

Occulting galaxies are often interacting and the redshift difference for two of our templates (UGC 3995 and 2MASX) are such that an ongoing interaction cannot be ruled out. It is very likely in the case of UGC 3995. The exact effect of gravitational interaction on the distribution of ISM is unknown but likely to be severe. Is there more dust now at higher radii? Do the shocks and tides move more dust into denser clouds in advance of the burst of star formation associated with gas-rich interactions? The way around this issue again is to observe more and bona fide occulting pairs, well separated in redshift. A selection from a spectroscopic survey such as the GAMA survey (Driver et al. 2009, 2011; Baldry et al. 2010) may well prove to be ideal.

5.5 Is the $A_V$ prior really an exponential?

One strong assumption so far has been that the distribution of extinction values is an exponential drop-off. However, Fig. 2 shows that none of the three distributions is fully described as an exponential. For example, in the more massive galaxies, the peak of the distribution is not close to $A_V = 0$ at all. We noted this in Holwerda & Keel (2013) for both the HST observations and the IFU ones. In the case of UGC 3995, one could contribute this to the ongoing interaction. Again the only way forward it to obtain more templates for the same host mass and infer the mean and variance of the distribution function.

Our assumption – and the MLCS2k2 authors’ – is that the shape of the distribution is an exponential. We now strongly suspect the true shape of the extinction distributions may be a lognormal one, similar to what Dalcanton et al. (in preparation) find for the extinction values in front of stars in M31 based on the PHAT project (Dalcanton et al. 2012) or an exponentially modified Gaussian. To illustrate, Fig. 9 shows fits to the $A_V$ distribution using the exponentially modified

5 CAVEATS

While the radial coverage and host mass range appears reasonably approximated by our three template host galaxies, there are a number of issues outstanding with just using three galaxies as templates. We will focus on the shortcomings of the occulting galaxy pairs presented as priors.
Figure 9. The shape of the $A_V$ distribution of the three occulting galaxies with enough overlap to generate a distribution. An exponentially modified Gaussian distribution was used with maximum likelihood to each distribution. The relevant parameters are listed in Table 3.

Table 3. The exponentially modified Gaussian fits to the $A_V$ distributions observed in occulting pairs.

<table>
<thead>
<tr>
<th>Pair</th>
<th>$m$</th>
<th>$\sigma$</th>
<th>$\lambda$</th>
</tr>
</thead>
<tbody>
<tr>
<td>2MASX</td>
<td>0.013</td>
<td>0.015</td>
<td>6.65</td>
</tr>
<tr>
<td>UGC3559</td>
<td>0.092</td>
<td>0.092</td>
<td>2.10</td>
</tr>
<tr>
<td>AM500</td>
<td>0.013</td>
<td>0.011</td>
<td>2.29</td>
</tr>
<tr>
<td>AM1316</td>
<td>1.41</td>
<td>0.50</td>
<td>70.60</td>
</tr>
</tbody>
</table>

Gaussian fit to each distribution using maximum likelihood (fit values in Table 3).

This strongly suggests that MLCS2k2 will now have to be redone using these templates instead of an exponential one.

6 APPLICATION TO SALT2

SALT2 handles the SN Ia colour and dust extinction in a combined parameter, and it may not be obvious how our approach has any easy application to this package. To illustrate, Fig. 10 shows the relation between the observed $A_V$ and the SALT2 colour parameter, $c$ (similar to fig. 9 in Sako et al. 2008). There is a linear relation between the two that follows:

$$c = 0.046 \times A_V + 0.012.$$  (5)

Thus, a change in the prior for $A_V$ in the case of MLCS2k2 would translate a change in the prior for SALT2’s assumption but in colour-space. To illustrate in Fig. 11, we plot the distribution of $c$ values found in SDSS-III for SN Ia, as well as the values of $A_V$ from MLCS2k2 translated using the above relation. The two distributions are – unsurprisingly – very similar. However, if we then translate the $A_V$ value corrected for the inclination of the disc (inferred from the axis ratio), the distribution of $c$ values changes significantly. In our view, this change in distribution after inclination correction is the original distribution one could adopt as the prior distribution of the parameter $c$ for SALT2 applications, similar to the one to be adopted for MLCS2k2 (since both of them are essentially colour priors).

7 DISCUSSION

The SN Ia community is well aware of the effects of host galaxy extinction may have on the results and have actively been searching for both ‘dust-free’ supernovae (i.e. hosted by early-type galaxies, see e.g. Suzuki et al. 2012) and pushed light-curve observations into the near-infrared to mitigate extinction issues.

This paper’s aim is an attempt to map out how the available observational information, and specifically the host galaxy information, can be further used to optimize the SN Ia distance measurement. As a first exercise, we applied the three available templates from occulting galaxies, naively matched to the SN Ia host through the stellar mass or galactocentric radius. Introducing new host extinction templates certainly has an effect but since the scatter in the distances is not markedly reduced, using just these templates would just skew the cosmological result, not improve it. Despite the fact that the $A_V$ values were attained using a more Milky Way-type relation (Kessler et al. 2009), we cannot rule out that none of the current...
scatter comes from host attenuation and all of it comes from intrinsic colour scatter in the SN Ia. But assuming much of the scatter is due to host attenuation, when matching host templates through projected radius rather than host mass, the scatter increases more than through mass-matching. This points, in our opinion, in which order one should match the occulting galaxy templates: first through mass and then through galactocentric radius.

Another key issue for SN Ia cosmological measurements will be if the extinction prior will change with redshift. Given that between $z=0$ and 1, specific star formation rates increased by a factor 3–10 (e.g. Noeske et al. 2007), one can expect a much more fractured dusty ISM (thanks to increased turbulence, as we speculated in (e.g. Noeske et al. 2007), one can expect a much more fractured dust distribution with redshifts – either due to ISM evolution in the host galaxies themselves. There is, at presence, no observed evolution in the colour distribution of observed SNe Ia (Conley et al. 2011; Betoule et al. 2014). However, if there is significant evolution in the $P(A_V)$ distribution with redshifts – either due to ISM evolution in the host galaxies or a significant change in the host galaxy population mix – it may change the relative luminosity distances as a function of redshift, the driver or dark energy cosmology.

We hope to have convinced the reader that the use of occulting galaxy pairs to obtain extinction distribution templates is a good way forward to help reduce the remaining uncertainties in SN Ia distance measurements. The eventual goal is to cover enough range in stellar mass, radial coverage, and inclinations to map the mean and variance of the typical extinction distribution in spiral galaxy discs. This will have many astrophysical applications but the use as a prior for cosmological SN Ia measurements was the principal driver for us to obtain HST observations of occulting pairs: GO-13695, Starlight Absorption Reduction through a Survey of Multiple Occulting Galaxies (STARMOG) P.I. B.W. Holwerda, a 150 orbit SNAPshot programme (Holwerda 2014), which is projected to complete a sample of 98 occulting pairs by 2017.

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Matching SN Ia dust extinction to a host 2397
