

## Outline

- 1 Scattering Polarization
- 2 Zeeman Effect
- 3 Hanle Effect

# Scattering Polarization

## Single Particle Scattering

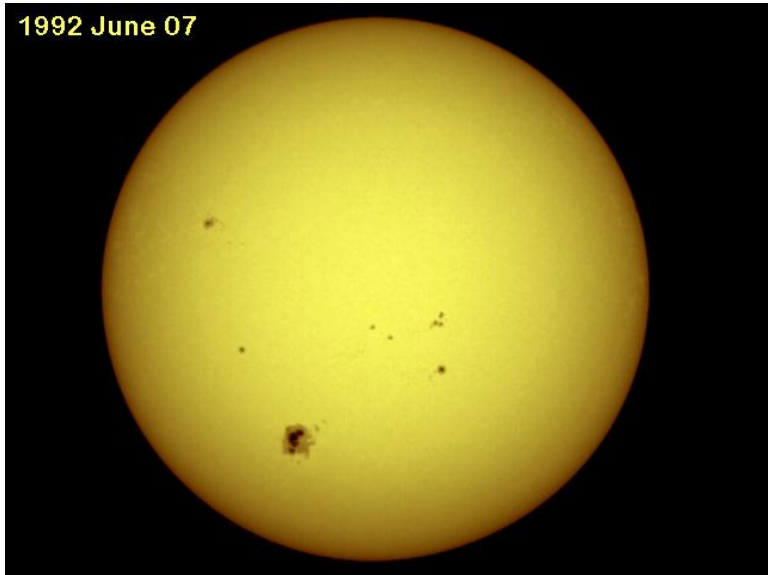
- light is absorbed and re-emitted
- if light has low enough energy, no energy transferred to electron, but photon changes direction  $\Rightarrow$  elastic scattering
- for high enough energy, photon transfers energy onto electron  $\Rightarrow$  inelastic (Compton) scattering
- Thomson scattering on free electrons
- Rayleigh scattering on bound electrons
- based on very basic physics, scattered light is linearly polarized

## Polarization as a Function of Scattering Angle

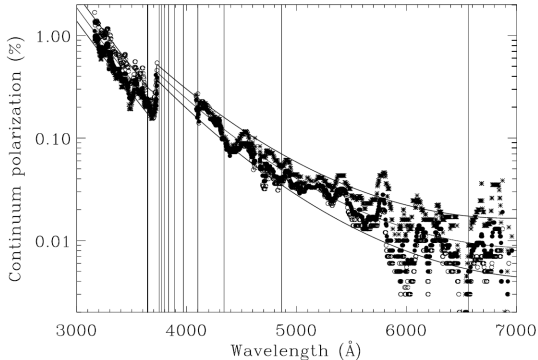
- same variation of polarization with scattering angle applies to Thomson and Rayleigh scattering
- scattering angle  $\theta$
- projection of amplitudes:
  - 1 for polarization direction perpendicular to scattering plane
  - $\cos \theta$  for linear polarization in scattering plane
- intensities = amplitudes squared
- ratio of  $+Q$  to  $-Q$  is  $\cos^2 \theta$  (to 1)
- total scattered intensity (unpolarized = averaged over all polarization states) proportional to  $\frac{1}{2} (1 + \cos^2 \theta)$

## Limb Darkening

1992 June 07



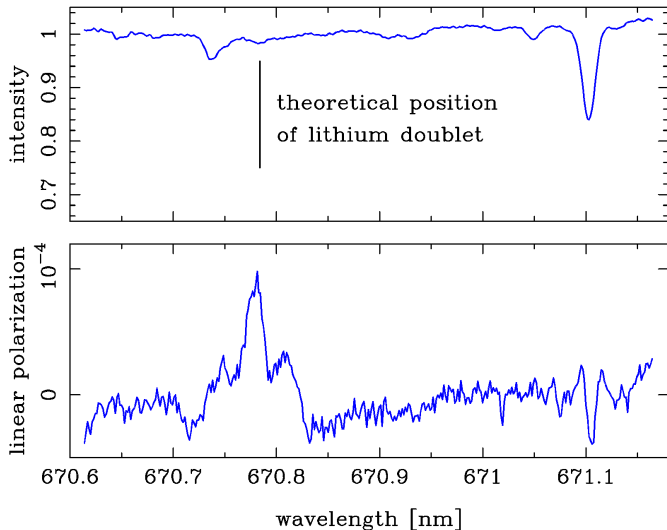
## Solar Continuum Scattering Polarization



Stenflo 2005

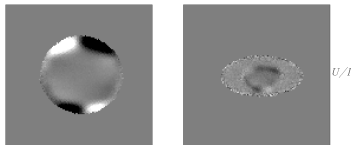
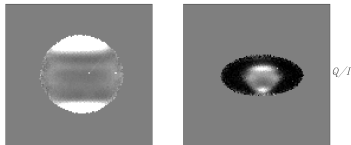
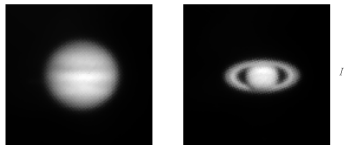
- due to anisotropy of the radiation field
- anisotropy due to limb darkening
- limb darkening due to decreasing temperature with height
- last scattering approximation without radiative transfer

## Solar Spectral Line Scattering Polarization



resonance lines exhibit "large" scattering polarization signals

## Jupiter and Saturn



(courtesy H.M.Schmid and D.Gisler)

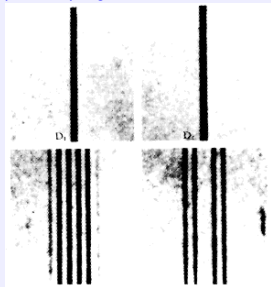
## Planetary Scattered Light

- Jupiter, Saturn show scattering polarization
- multiple scattering changes polarization as compared to single scattering
- much depends on cloud height
- can be used to study extrasolar planetary systems
- ExPo instrument development at UU

# Zeeman Effect



[photos.aip.org/](http://photos.aip.org/)



## Splitting/Polarization of Spectral Lines

- discovered in 1896 by Dutch physicist Pieter Zeeman
- different spectral lines show different splitting patterns
- splitting proportional to magnetic field
- split components are polarized
- *normal Zeeman effect* with 3 components explained by H.A.Lorentz using classical physics
- splitting of sodium D doublet could not be explained by classical physics (*anomalous Zeeman effect*)
- quantum theory and electron's intrinsic spin led to satisfactory explanation



## Quantum-Mechanical Hamiltonian

- classical interaction of magnetic dipol moment  $\vec{\mu}$  and magnetic field given by magnetic potential energy

$$U = -\vec{\mu} \cdot \vec{B}$$

$\vec{\mu}$  the magnetic moment and  $\vec{B}$  the magnetic field vector

- magnetic moment of electron due to orbit and spin
- Hamiltonian for quantum mechanics

$$H = H_0 + H_1 = H_0 + \frac{e}{2mc} (\vec{L} + 2\vec{S}) \cdot \vec{B}$$

$H_0$  Hamiltonian of atom without magnetic field

$H_1$  Hamiltonian component due to magnetic field

$e$  charge of electron

$m$  electron rest mass

$\vec{L}$  the orbital angular momentum operator

$\vec{S}$  the spin operator

## Energy States in a Magnetic Field

- energy state  $\langle E_{NLSJ} |$  characterized by
  - main quantum number  $N$  of energy state
  - $L(L + 1)$ , the eigenvalue of  $\vec{L}^2$
  - $S(S + 1)$ , the eigenvalue of  $\vec{S}^2$
  - $J(J + 1)$ , the eigenvalue of  $\vec{J}^2$ ,  
 $\vec{J} = \vec{L} + \vec{S}$  being the total angular momentum
  - $M$ , the eigenvalue of  $J_z$  in the state  $\langle NLSJM |$
- for the magnetic field in the z-direction, the change in energy is given by

$$\Delta E_{NLSJ}(M) = \langle NLSJM | H_1 | NLSJM \rangle$$

## The Landé g Factor

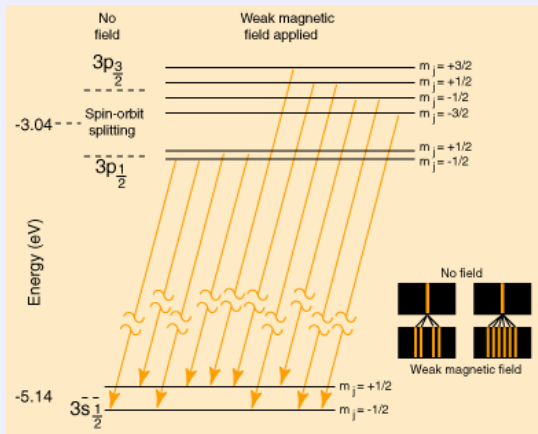
- based on pure mathematics (group theory, Wigner-Eckart theorem), one obtains

$$\Delta E_{NLSJ}(M) = \mu_0 g_L B M$$

with  $\mu_0 = \frac{e\hbar}{2m}$  the Bohr magneton, and  $g_L$  the Landé g-factor

- in LS coupling where  $B$  sufficiently small compared to spin-orbit splitting field

$$g_L = 1 + \frac{J(J+1) + L(L+1) - S(S+1)}{2J(J+1)}$$



[hyperphysics.phy-astr.gsu.edu/hbase/quantum/sodzee.html](http://hyperphysics.phy-astr.gsu.edu/hbase/quantum/sodzee.html)

## Spectral Lines - Transitions between Energy States

- spectral lines are due to transitions between energy states:

lower level with  $2J_l + 1$  sublevels  $M_l$

upper level with  $2J_u + 1$  sublevels  $M_u$

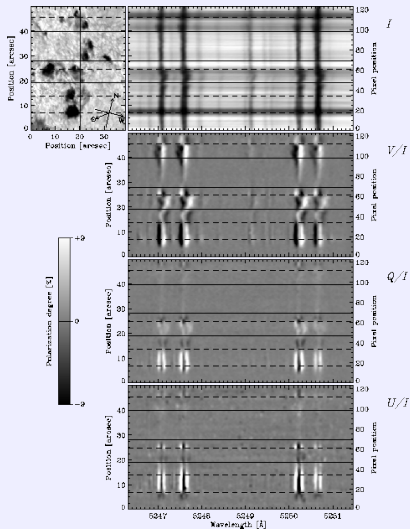
- not all transitions occur

## Selection rule

- not all transitions between two levels are allowed
- assuming dipole radiation, quantum mechanics gives us the *selection rules*:
  - $L_u - L_l = \Delta L = \pm 1$
  - $M_u - M_l = \Delta M = 0, \pm 1$
  - $M_u = 0$  to  $M_l = 0$  is forbidden for  $J_u - J_l = 0$
- total angular momentum conservation: photon always carries  $J_{\text{photon}} = 1$
- *normal Zeeman effect*: line splits into three components because
  - Landé g-factors of upper and lower levels are identical
  - $J_u = 1$  to  $J_l = 0$  transition
- *anomalous Zeeman effect* in all other cases

## Effective Landé Factor and Polarized Components

- each component can be assigned an effective Landé g-factor, corresponding to how much the component shifts in wavelength for a given field strength
- components are also grouped according to the linear polarization direction for a magnetic field perpendicular to the line of sight
  - $\pi$  components are polarized parallel to the magnetic field (**p**i for *parallel*)
  - $\sigma$  components are polarized perpendicular to the magnetic field (**s**igma for German *senkrecht*)
- for a field parallel to the line of sight, the  $\pi$ -components are not visible, and the  $\sigma$  components are circularly polarized

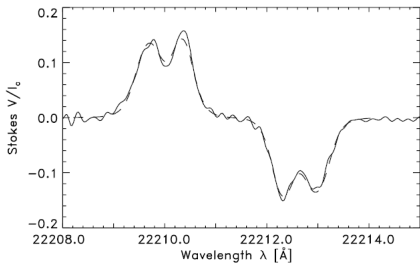
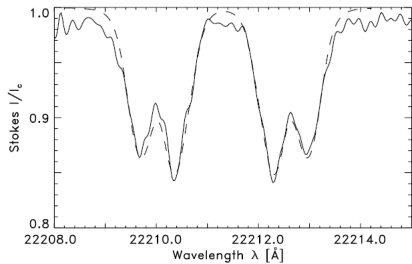
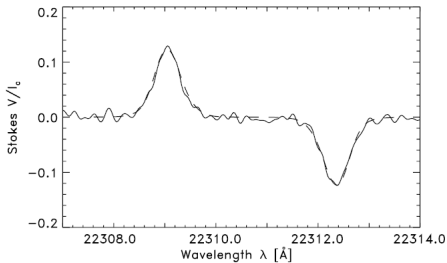
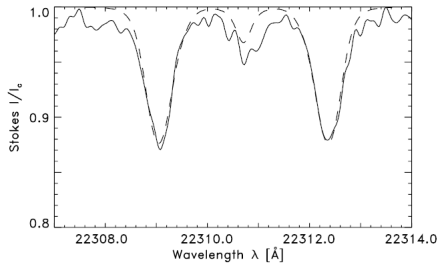


Bernasconi et al. 1998

## Zeeman Effect in Solar Physics

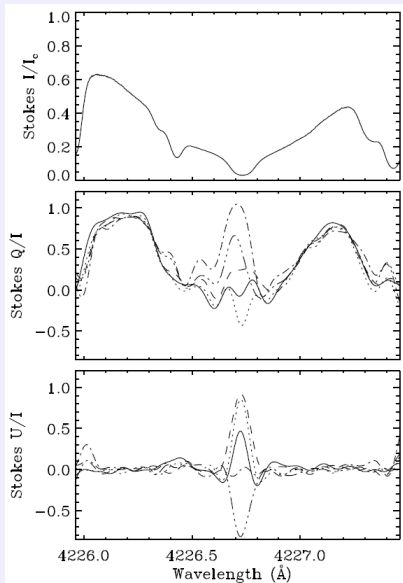
- discovered in sunspots by G.E.Hale in 1908
- splitting small except for in sunspots
- much of intensity profile due to non-magnetic area  $\Rightarrow$  filling factor
- a lot of strong fields outside of sunspots
- full Stokes polarization measurements are key to determine solar magnetic fields
- 180 degree ambiguity

# Fully Split Titanium Lines at 2.2 $\mu\text{m}$





# Hanle Effect



## Depolarization and Rotation

- scattering polarization modified by magnetic field
- precession around magnetic field depolarizes and rotates polarization
- sensitive  $\sim 10^3$  times smaller field strengths than Zeeman effect
- measurable effects even for isotropic field vector orientations