Astronomical Observing Techniques

Lecture 13:
Twinkle, twinkle little star …
No more!

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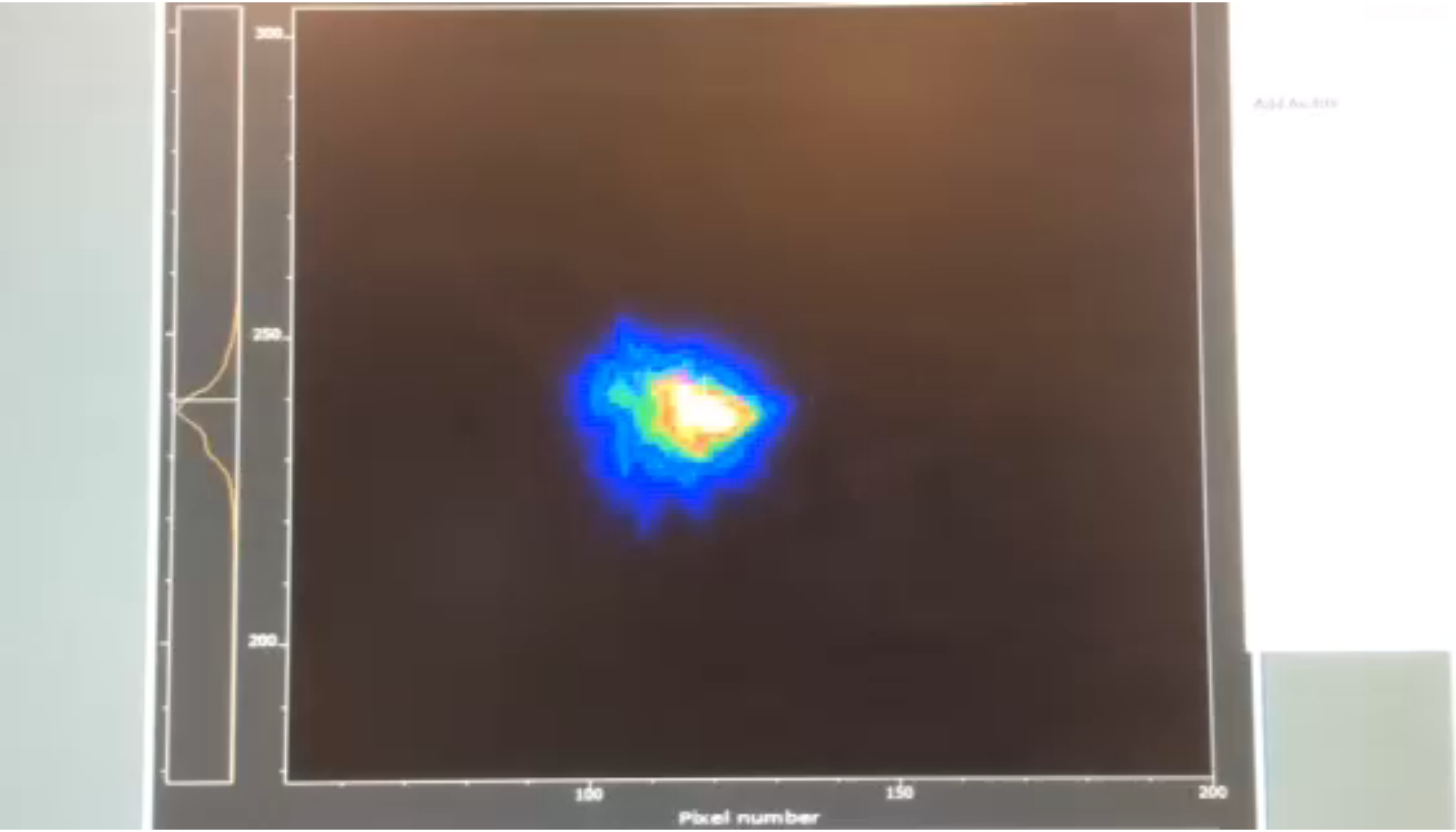
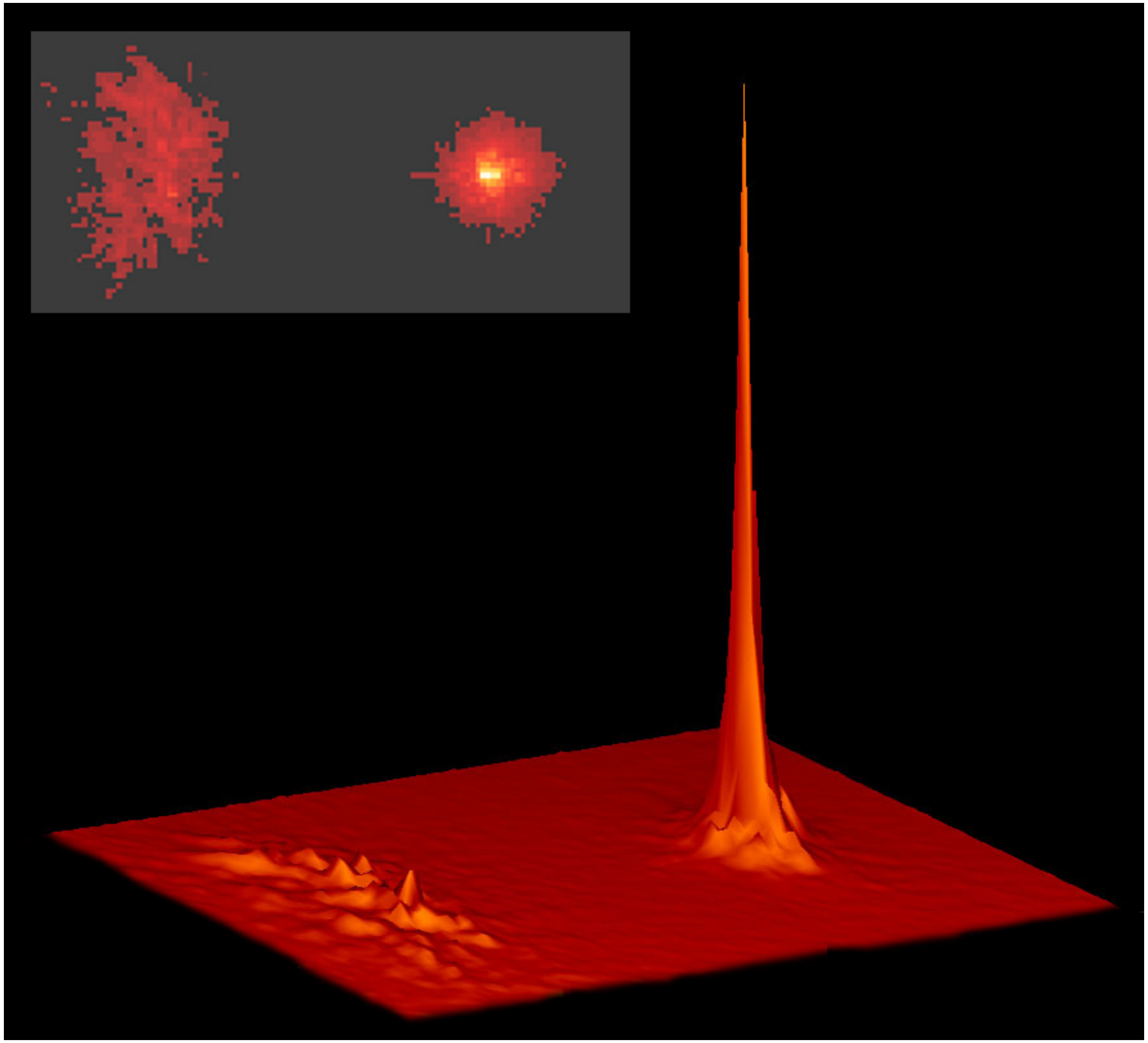
Overview

1. The Power of Adaptive Optics
2. Atmospheric Turbulence
3. Basic Principle
4. Key Components
5. Error Terms
6. Laser Guide Stars
7. Adaptive Optics Operations Modes

ExPo Adaptive Optics at WHT

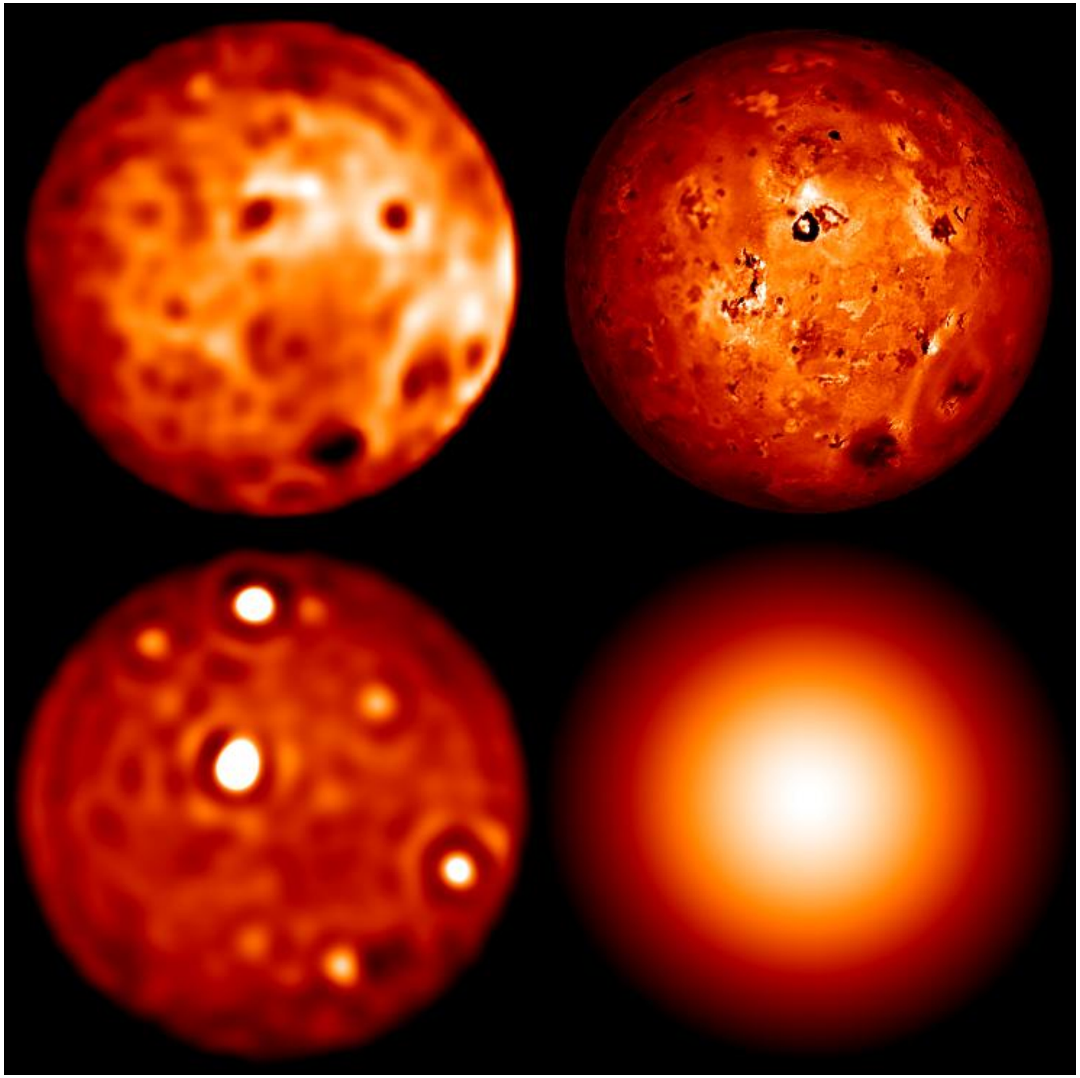


Star with and without AO



cfao.ucolick.org/pgallery/stellar.php

Io with and without Keck AO



Io image taken with Keck adaptive optics; K-band, **2.2micron**

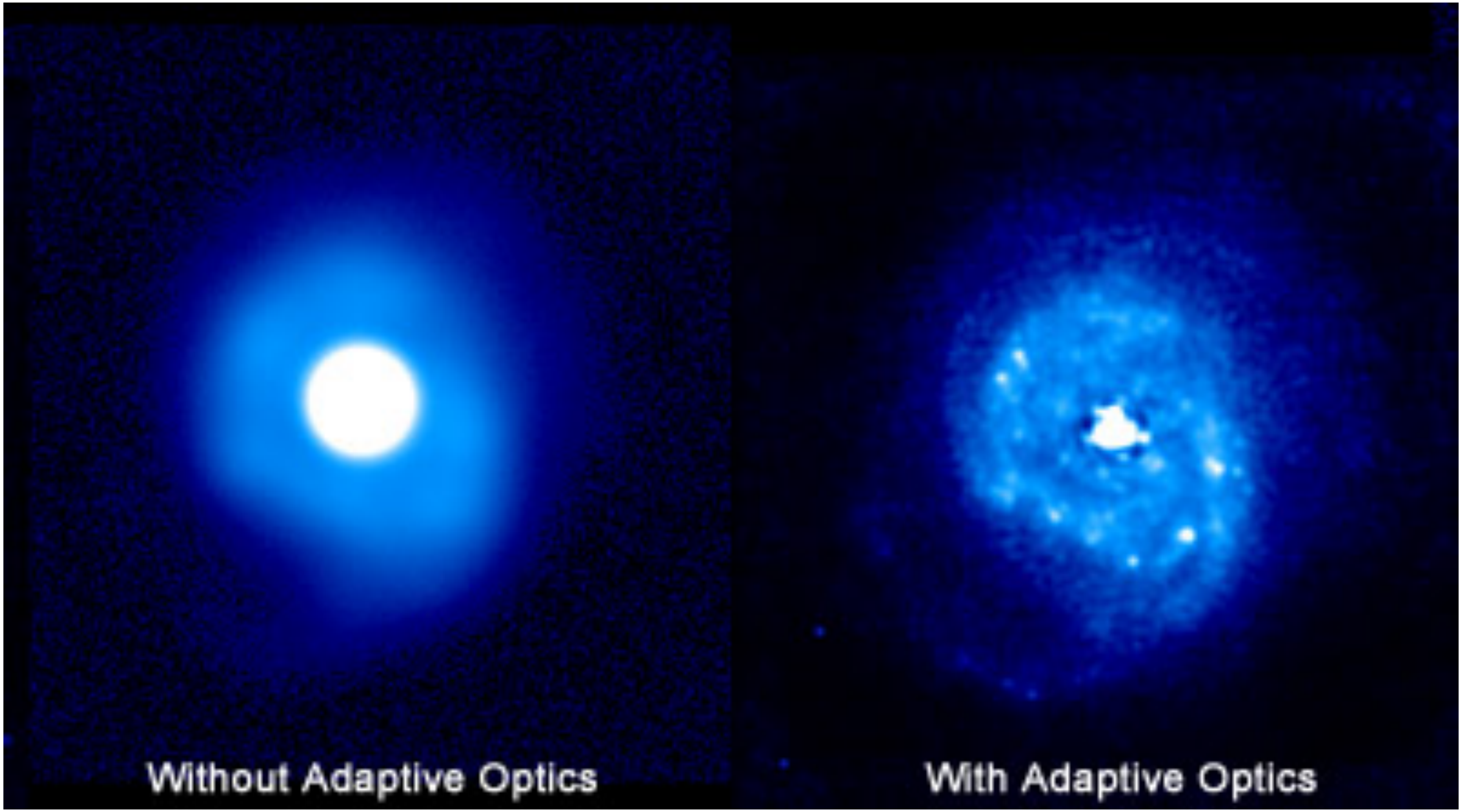
Io image based on visible light taken with **Galileo** spacecraft orbiter.

Io image taken with Keck adaptive optics; L-band, **3.5micron**

Io image taken **without** Keck adaptive optics.

cfao.ucolick.org/pgallery/io.php

NGC 7469



cfao.ucolick.org/ao/why.php

Seeing: r_0, τ_0, θ_0 (from Lecture 2)

- Fried parameter r_0: average turbulent scale over which RMS optical phase distortion is 1 rad

$$r_0(\lambda) = 0.185\lambda^{6/5}\left[\int_0^\infty C_n^2(z)dz\right]^{-3/5}$$

- r_0 increases as $\lambda^{6/5}$
- Seeing $\Delta\theta$ at good sites at 0.5µm: 10 - 30 cm

$$\Delta\theta = \frac{\lambda}{r_0} \sim \lambda^{-1/5}$$

- atmospheric coherence (or Greenwood delay) time: maximum time delay for RMS wavefront error to be < 1 rad ($\bar{v}$ is mean propagation velocity)

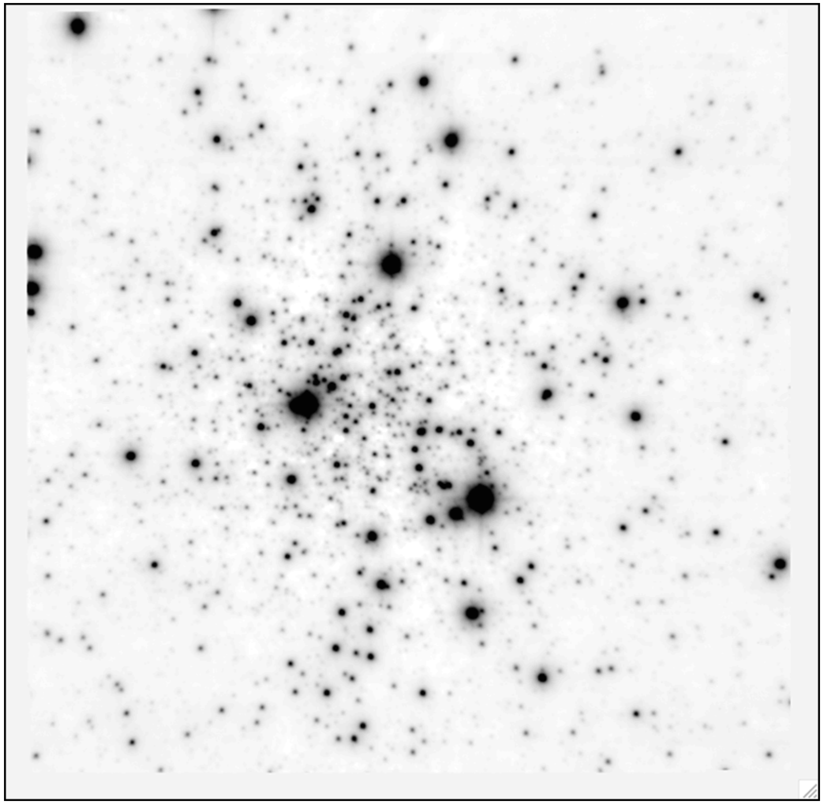
$$\tau_0 = 0.314\frac{r_0}{\bar{v}}$$

- Isoplanatic angle θ_0: angle over which RMS wavefront error is smaller than 1 rad

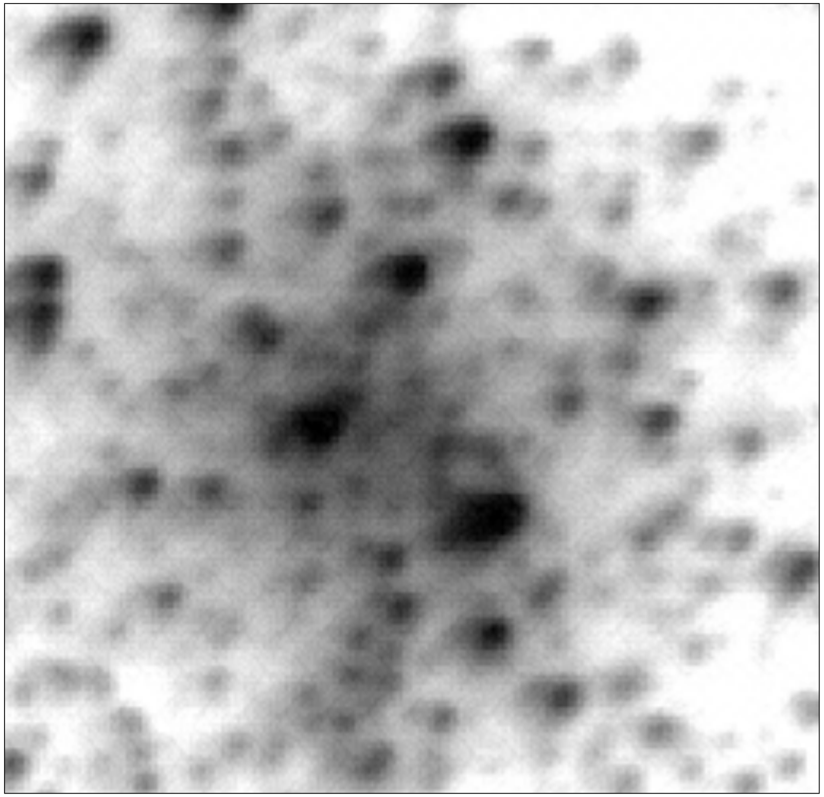
$$\theta_0 = 0.314\cos\zeta\frac{r_0}{\bar{h}}$$

Resolution & Sensitivity Improvement

1. Angular resolution: $\theta = \frac{\lambda}{r_0} \rightarrow \theta = \frac{\lambda}{D} \qquad \Rightarrow \text{gain} = \frac{D}{r_0}$
2. Point source sensitivity: $S/N \sim \mathrm{D}^2 \qquad \Rightarrow \text{gain in } t_{\text{int}} \sim \frac{1}{D^4}$

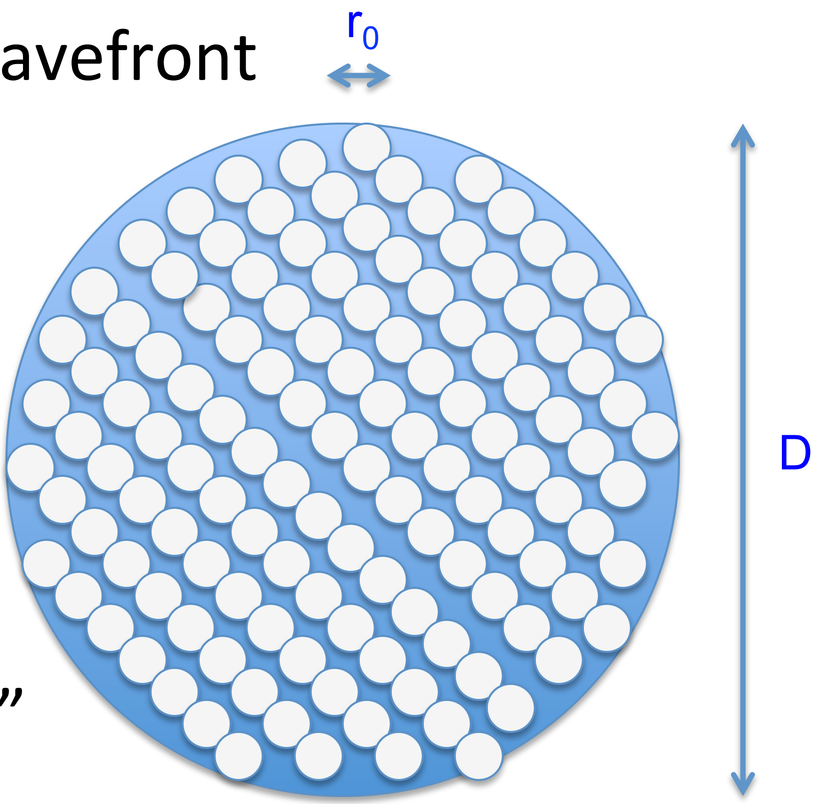


PHARO LGS Ks image
500s integ., 40" FOV, 150 mas FWHM



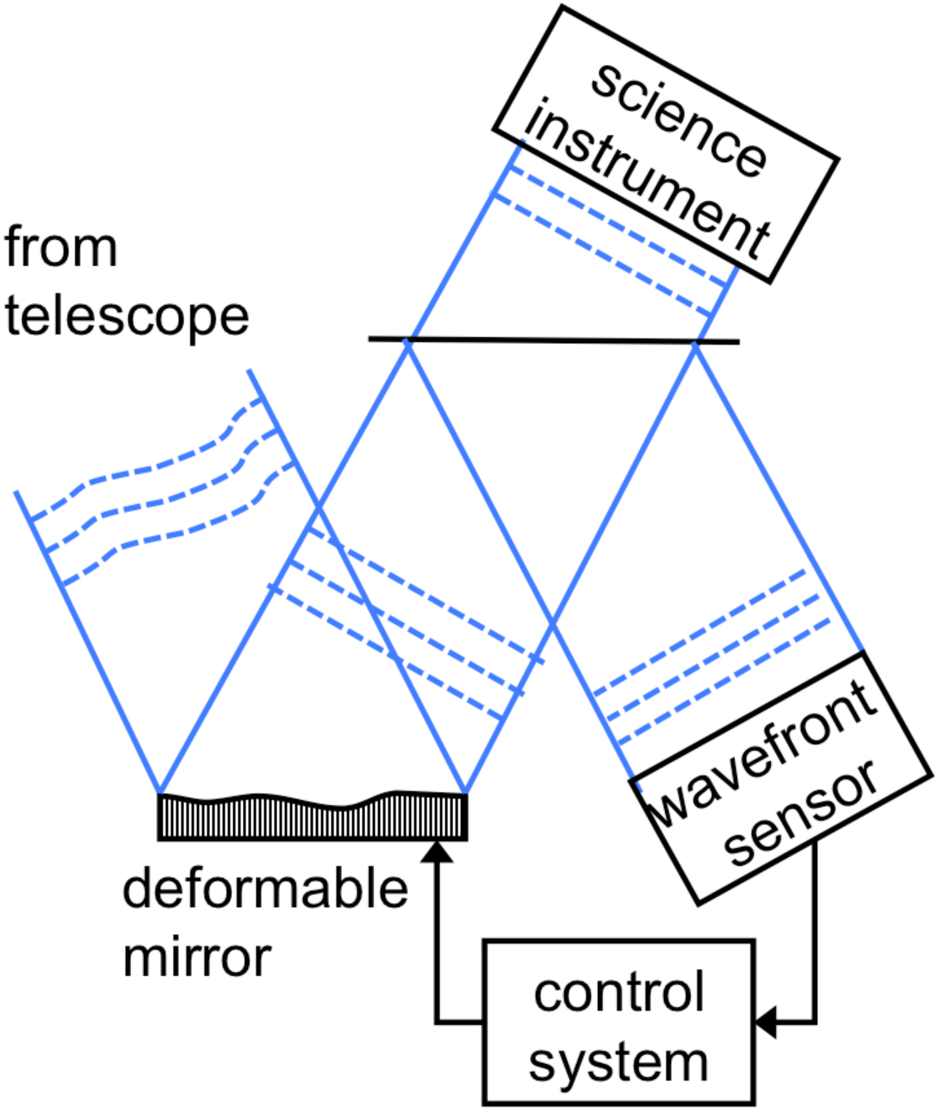
WIRO H image
Kobulnicky et al. 2005, AJ 129, 239-250

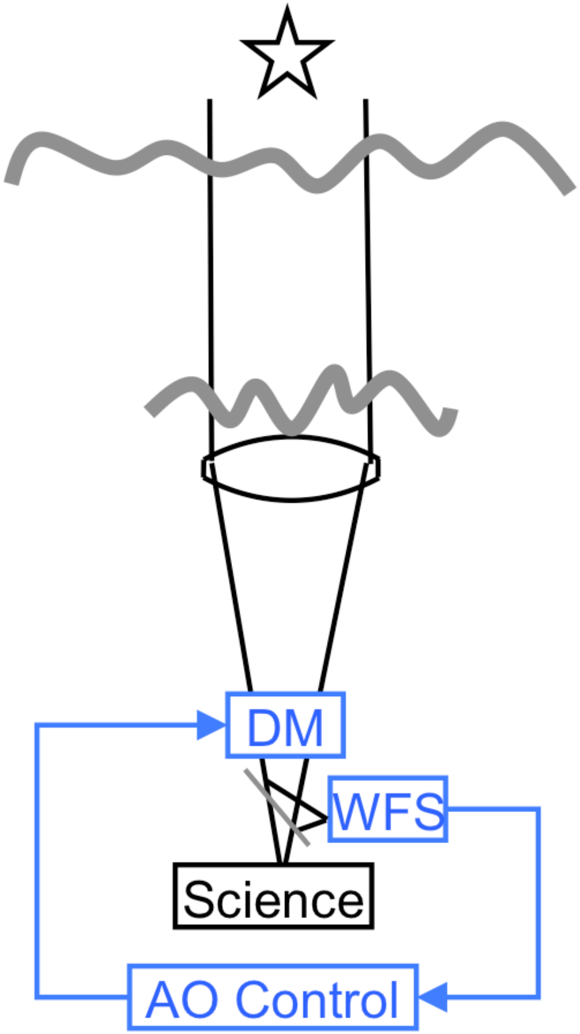
Adaptive Optics Principle



- Maximum scale of tolerated wavefront deformation is r_0 → subdivide telescope into apertures with diameter r_0
- Measure wavefront deformation
- Correct wavefront deformation by "bending back" patches of size r_0
- Number of subapertures is $(D/r_0)^2$ at observing wavelength → requires hundreds to thousands of actuators for large telescopes

Adaptive Optics Scheme (SCAO)



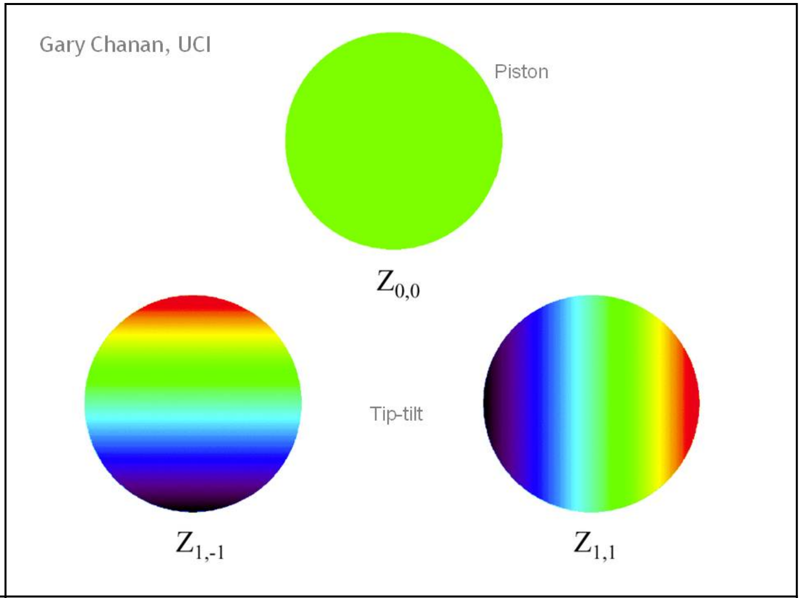

Single Conjugate AO

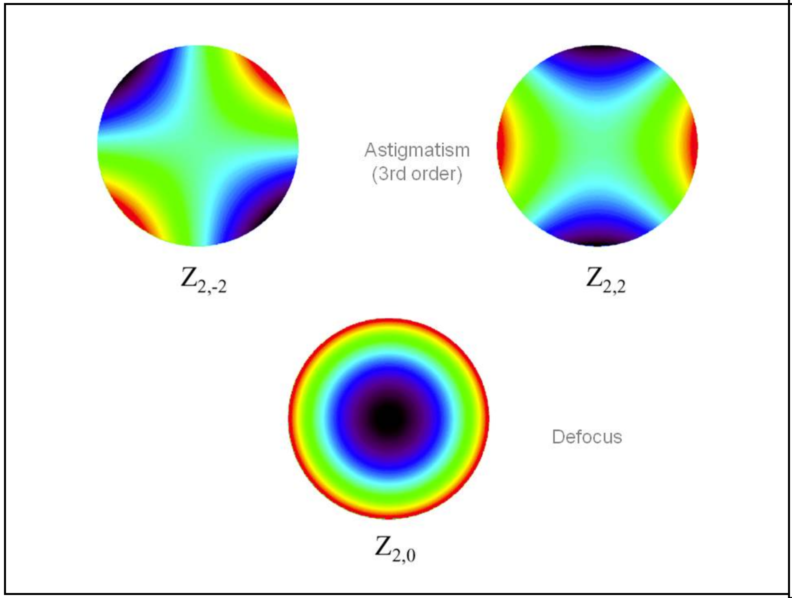
Wavefront Description: Zernike Polynomials

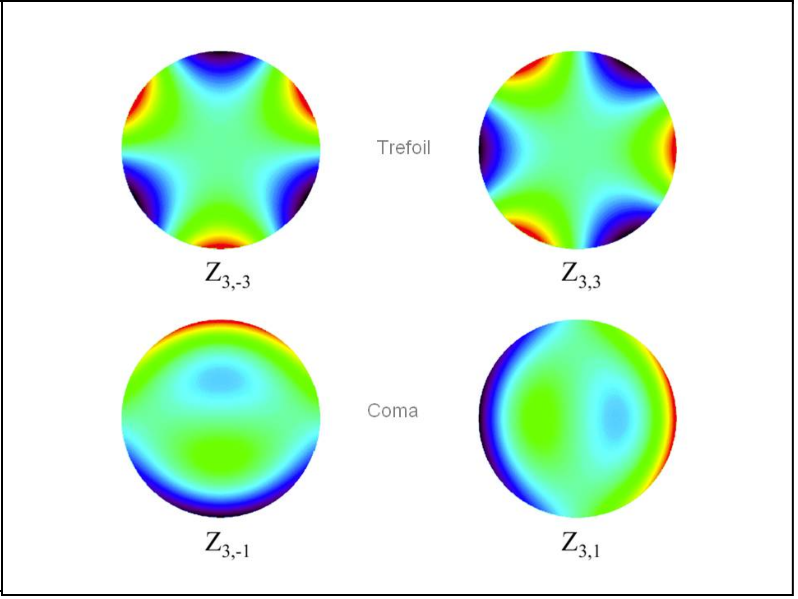
Expansion into a series of orthogonal terms:

$$\varphi(r,\theta) = \Sigma\, a_{m,n} Z_{m,n}(r,\theta)$$

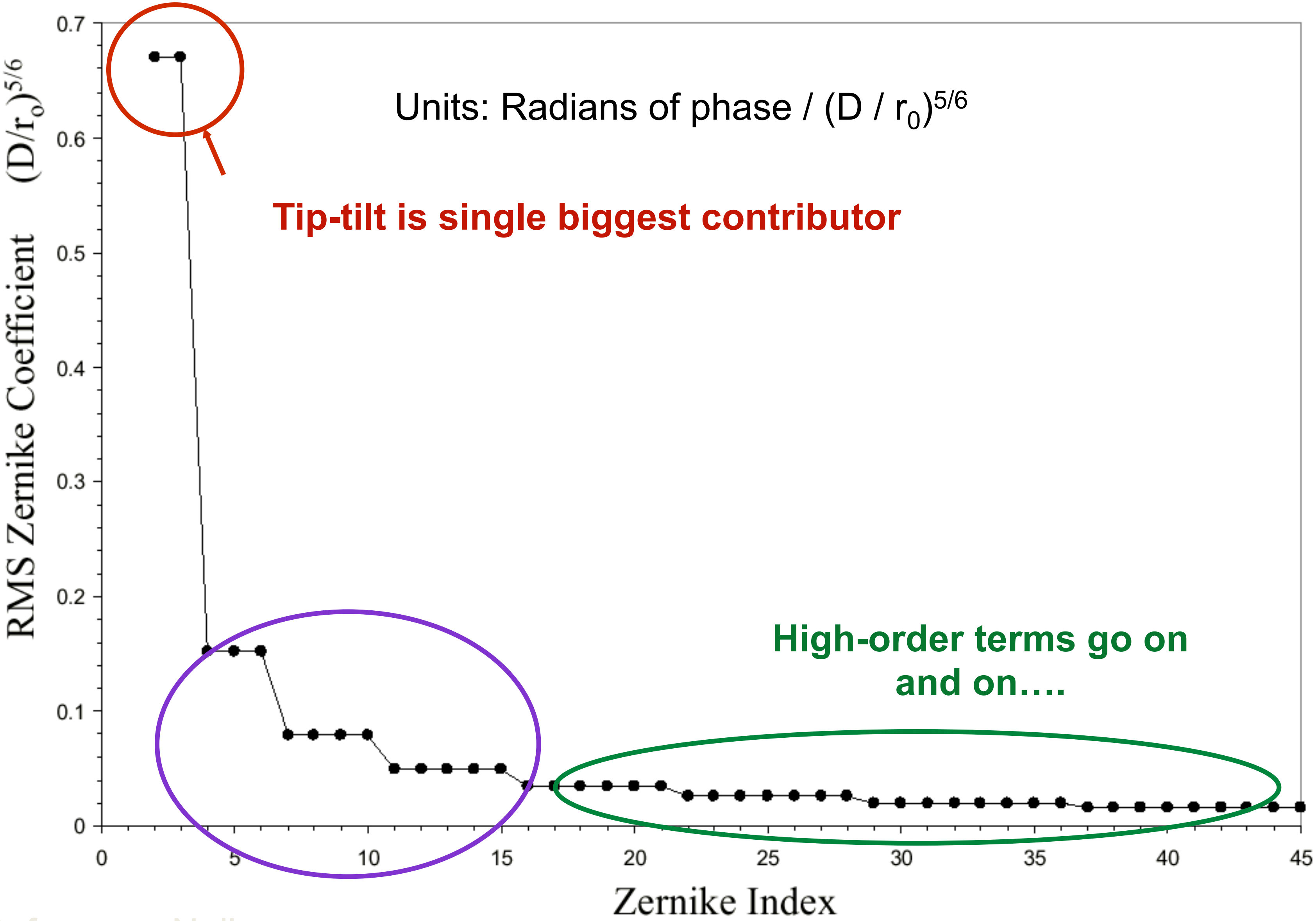
$Z_{0,0} = 1$	piston
$Z_{1,-1} = 2\,r \sin\theta$	tip/tilt
$Z_{1,1} = 2\,r \cos\theta$	
$Z_{2,-2} = \sqrt{6}\, r^2 \sin 2\theta$	astigmatism
$Z_{2,0} = \sqrt{3}\,(2r^2 - 1)$	focus
$Z_{2,2} = \sqrt{6}\, r^2 \cos 2\theta$	astigmatism


Gary Chanan, UCI





Zernike Amplitudes for Kolmogorov Turbulence



Units: Radians of phase / $(D / r_0)^{5/6}$

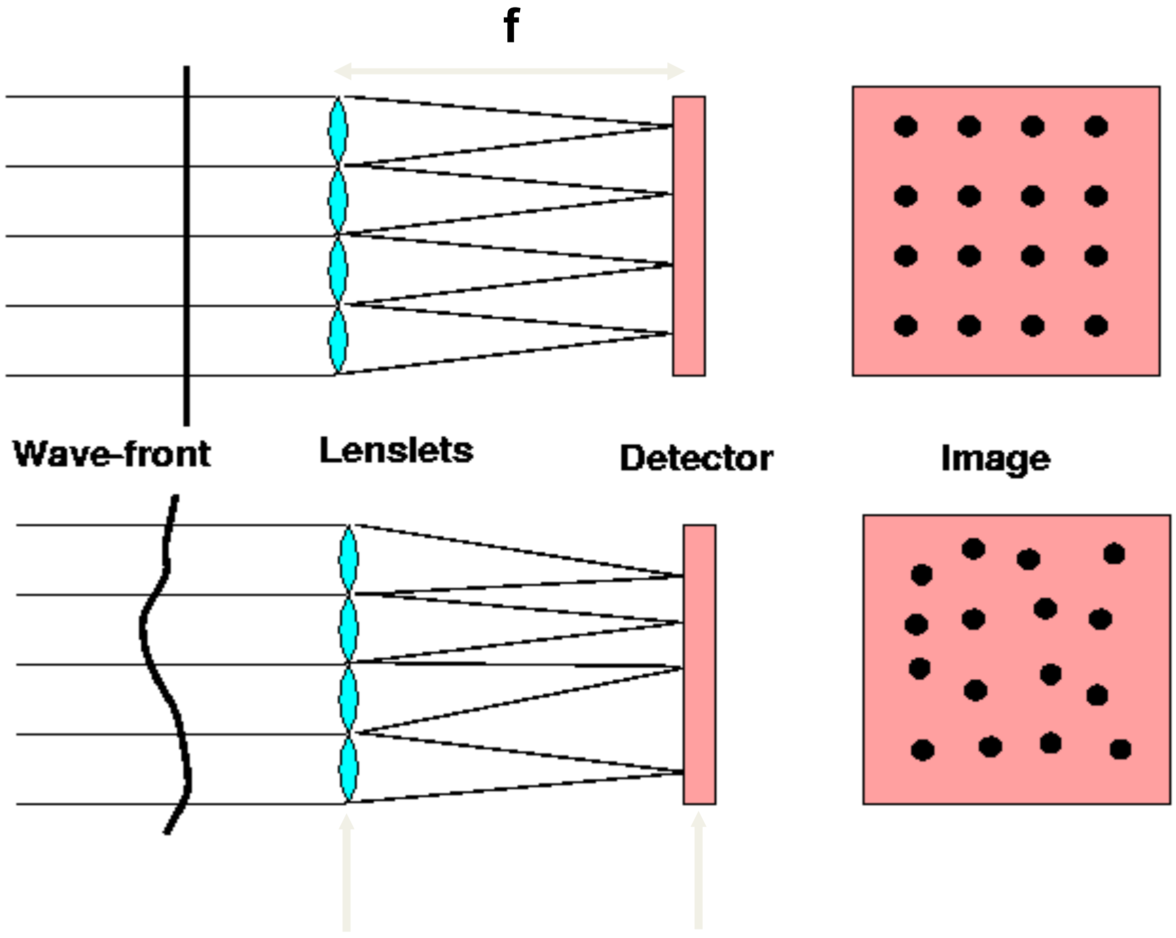
Tip-tilt is single biggest contributor

High-order terms go on and on….

Reference: Noll

Wavefront Sensors – Shack Hartmann

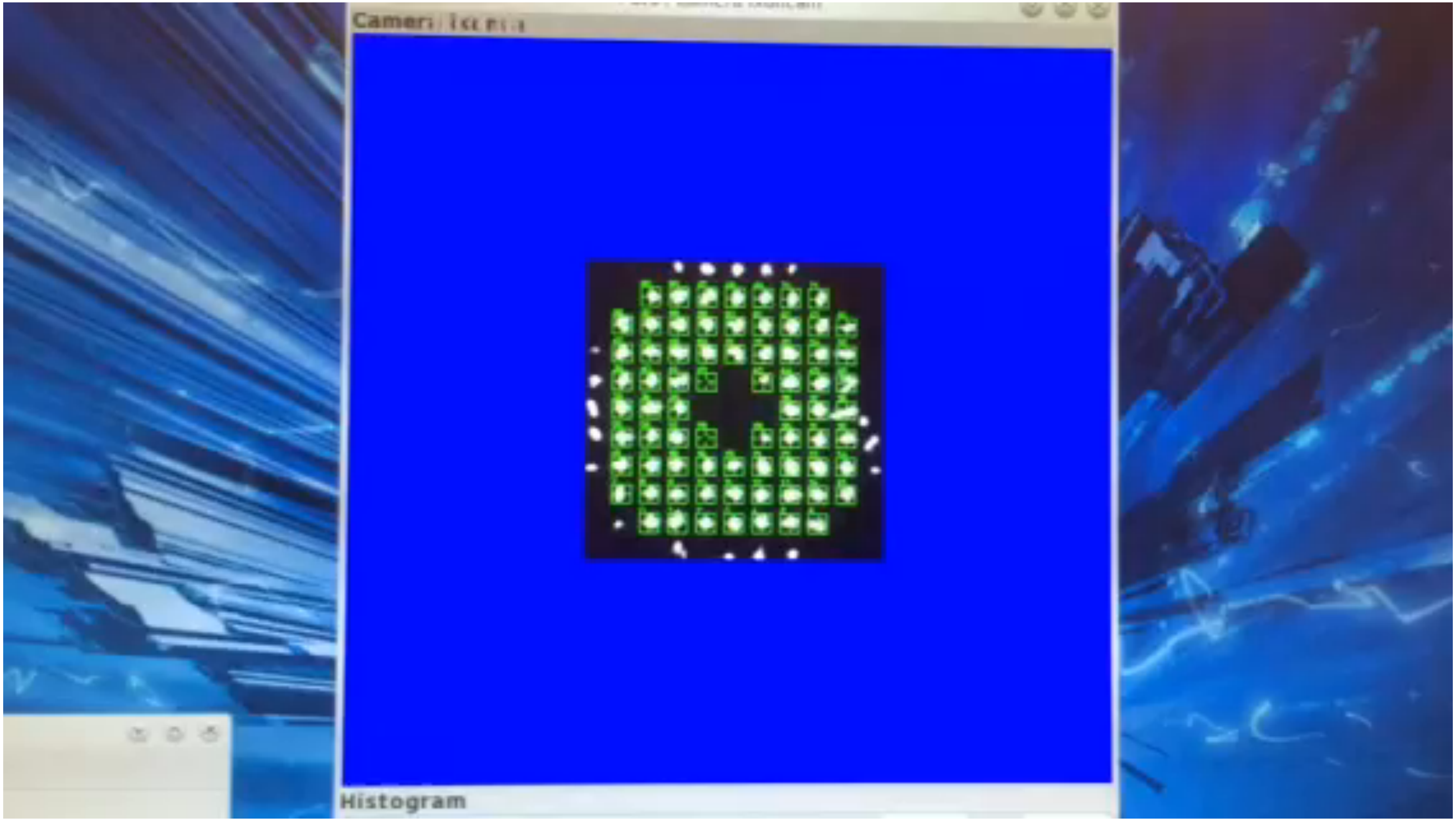
Most common principle is the Shack Hartmann wavefront sensor measuring sub-aperture tilts:



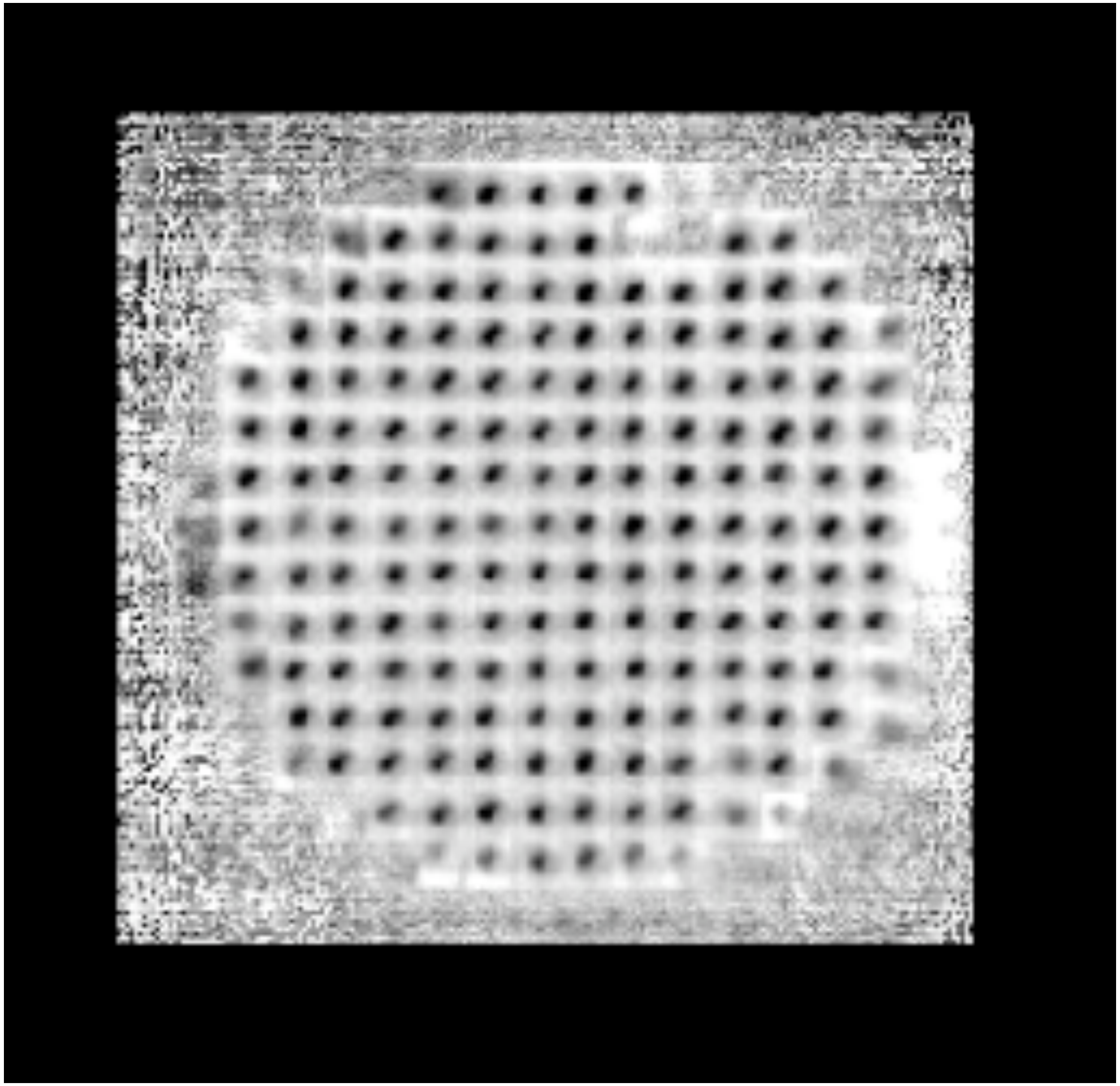
Pupil plane

Image plane

ExPo Wavefront Sensor at WHT



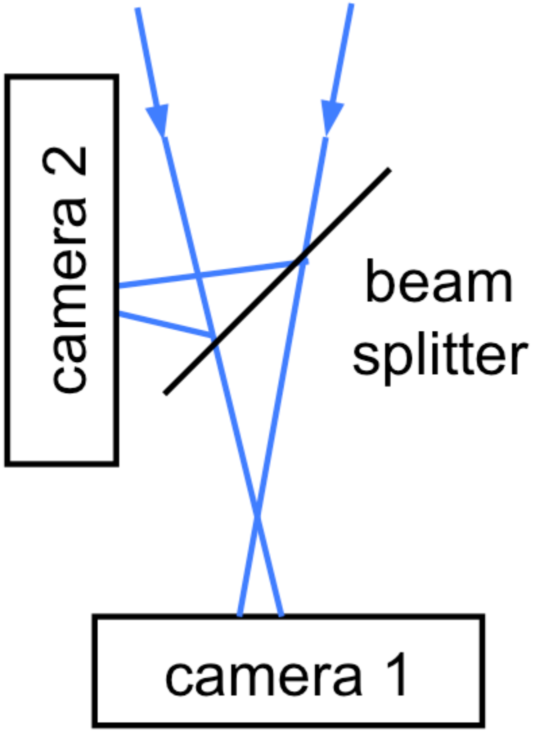
Wavefront Sensing



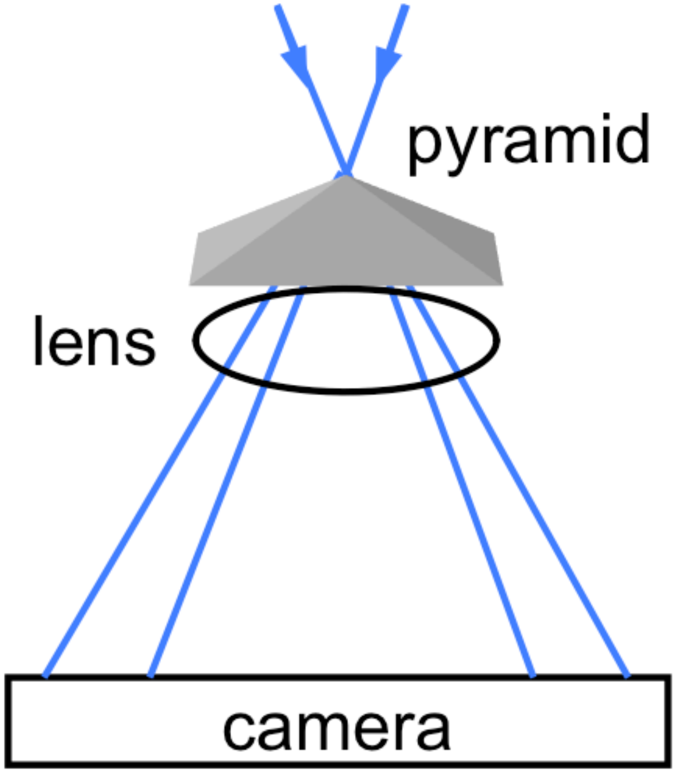
Sunspot wavefront sensor images at 955Hz

Curvature, Pyramid Wavefront Sensors

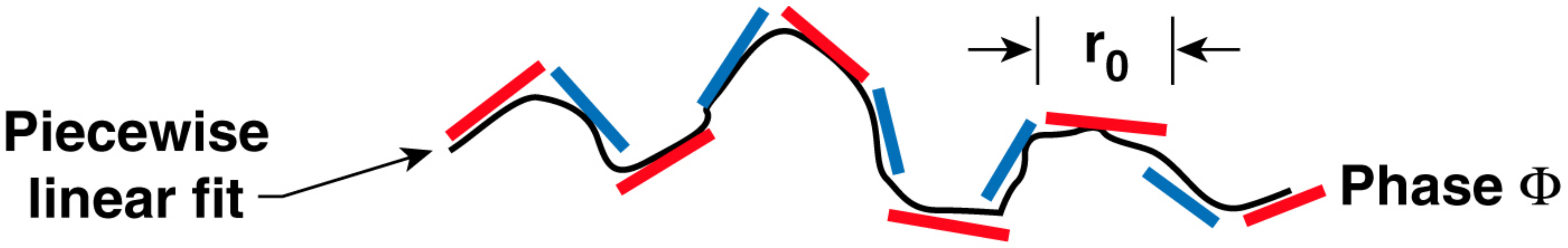
Curvature Wavefront Sensor



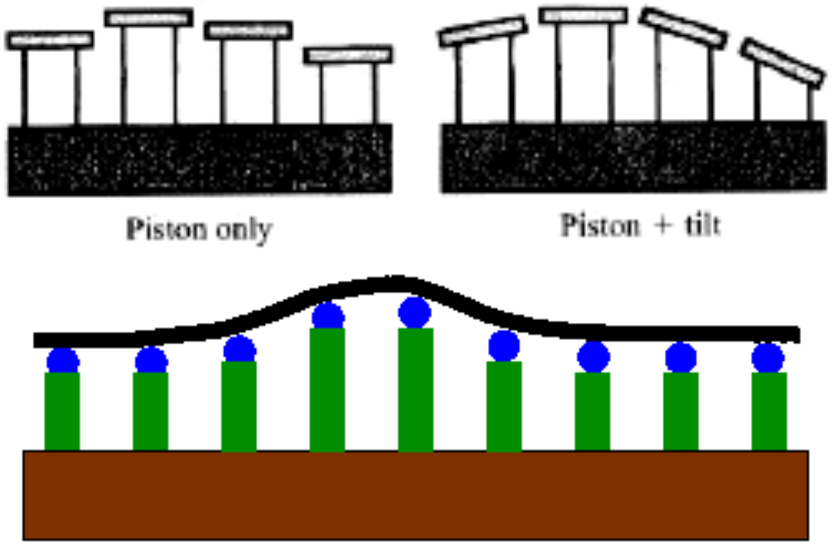
Pyramid Wavefront Sensor



Deformable Mirrors (DM)

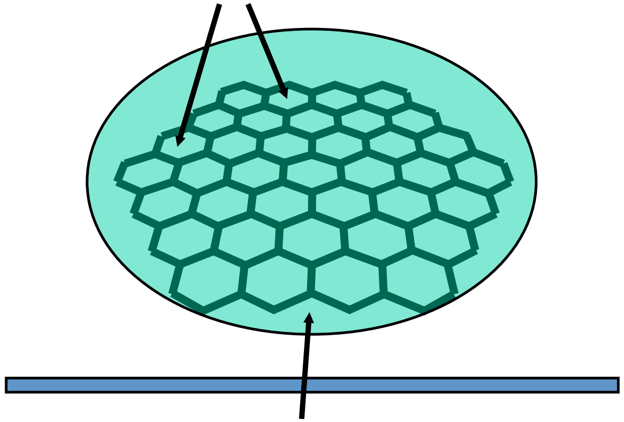


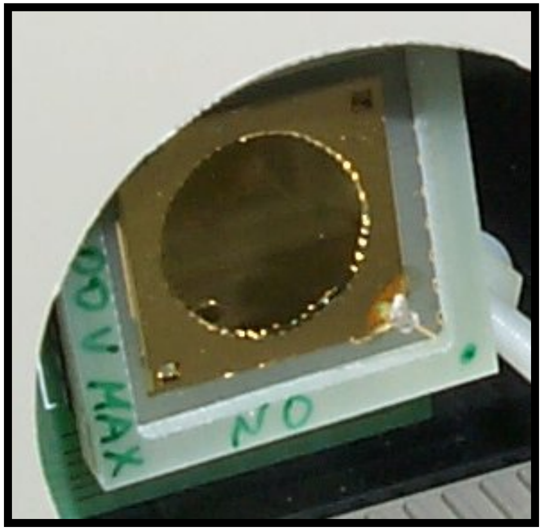
- Fit mirror surface to wavefront
- r_0 sets number of degrees of freedom
- segmented mirrors rarely used anymore
- mostly continuous face-sheet mirrors

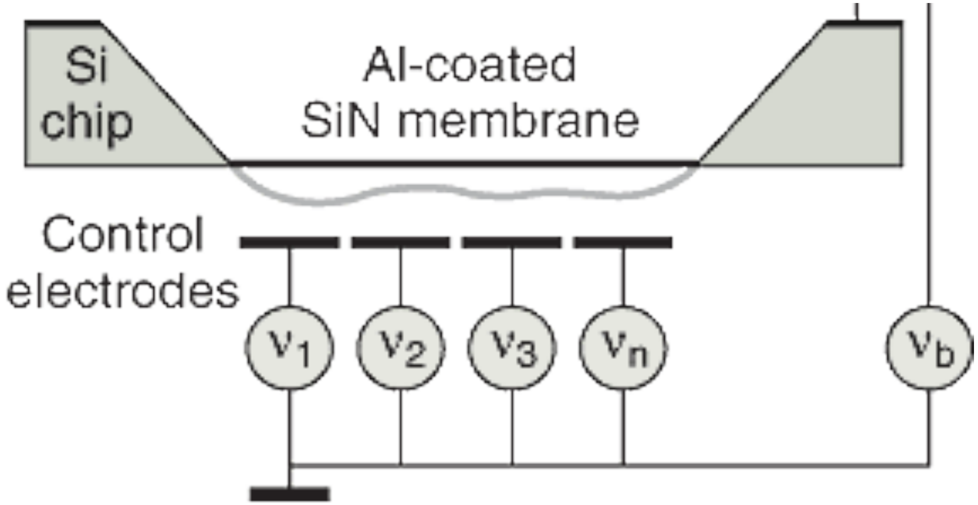


Membrane Deformable Mirror

- micromachined deformable mirror (OKOtech/Flexible Optics) with 37 actuators
- 600-nm thick, 15-mm diameter silicon nitride membrane
- electrostatic actuators

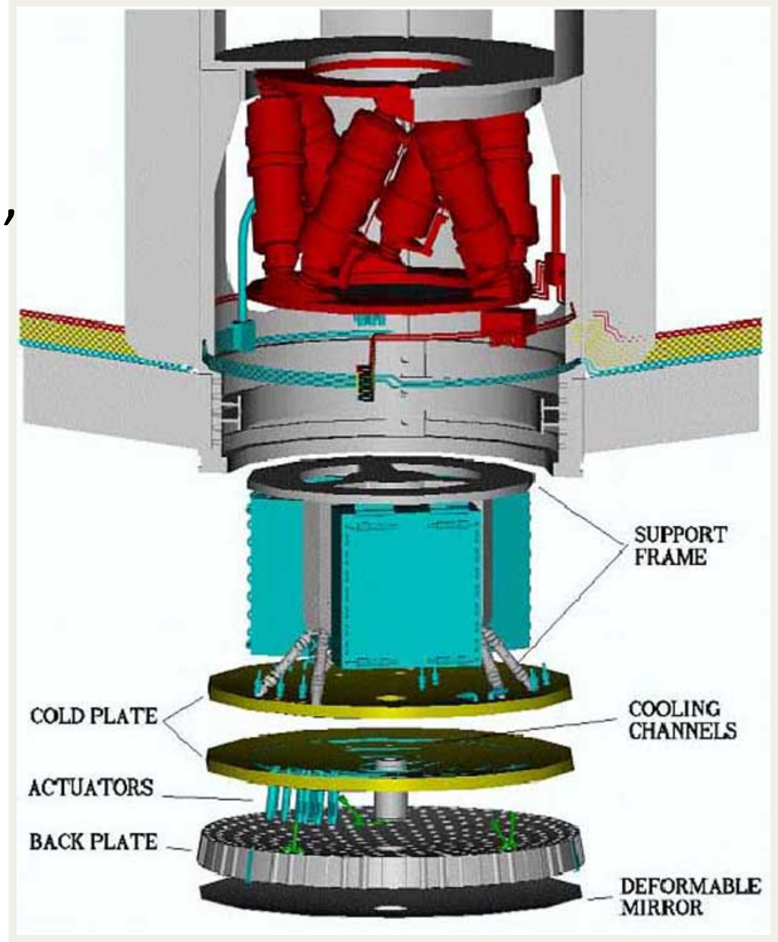


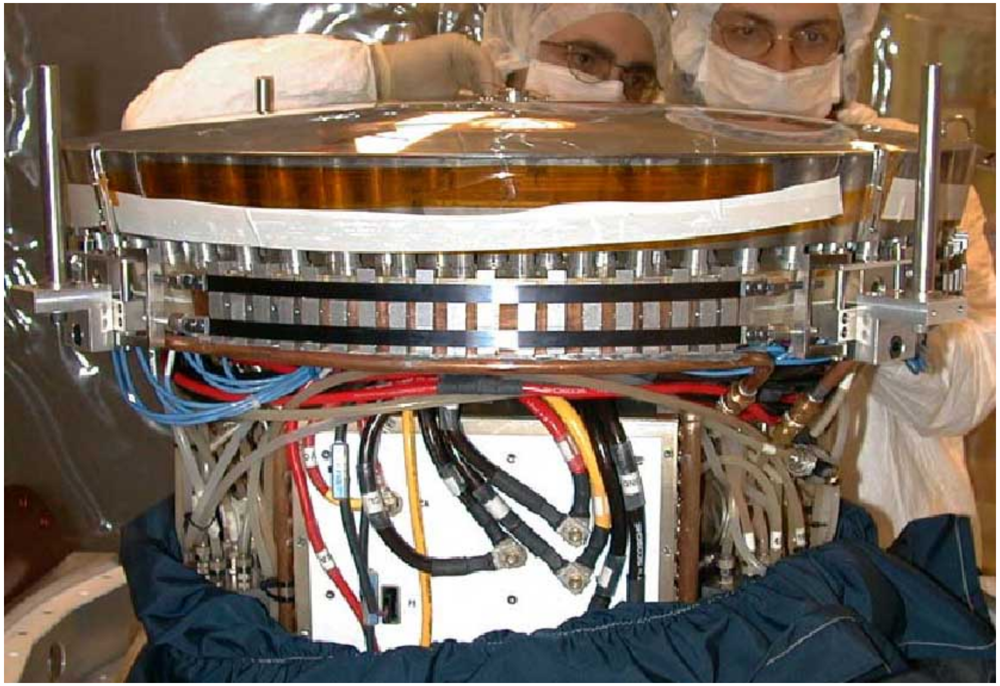




Adaptive Secondary Mirrors

- DM part of telescope → adaptive secondary mirrors
- no additional optics → lower emission, higher throughput
- more difficult to build, control, and handle



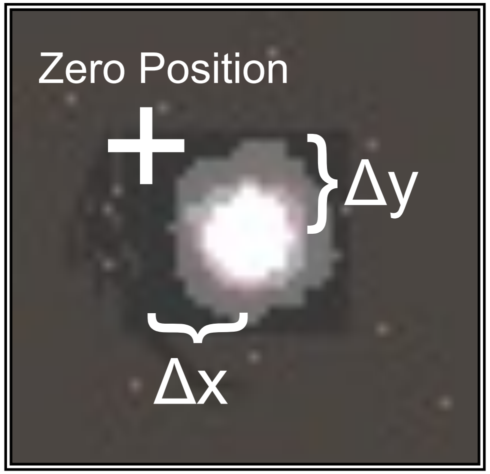


DM for MMT Upgrade

Shack-Hartmann Wavefront Analysis

- centroid (center of gravity) calculation on each subaperture:

$$\Delta x = \frac{\sum_{i=i_{\min}}^{i_{\max}} \sum_{j=j_{\min}}^{j_{\max}} I(i,j)\cdot i}{\sum_{i=i_{\min}}^{i_{\max}} \sum_{j=j_{\min}}^{j_{\max}} I(i,j)} \qquad \Delta y = \frac{\sum_{i=i_{\min}}^{i_{\max}} \sum_{j=j_{\min}}^{j_{\max}} I(i,j)\cdot j}{\sum_{i=i_{\min}}^{i_{\max}} \sum_{j=j_{\min}}^{j_{\max}} I(i,j)}$$



Influence Matrix

- slope of mirror surface *and* Shack-Hartmann star positions are proportional to actuator position
- linear relationship between actuator *a* and star position *c*:

$$c_n = \sum_{k=1}^{N} a_k b_{nk}$$ (For a single spot — x-offset or y-offset)

- combine equations for each spot position *n* into matrix equation:

$$C = BA$$

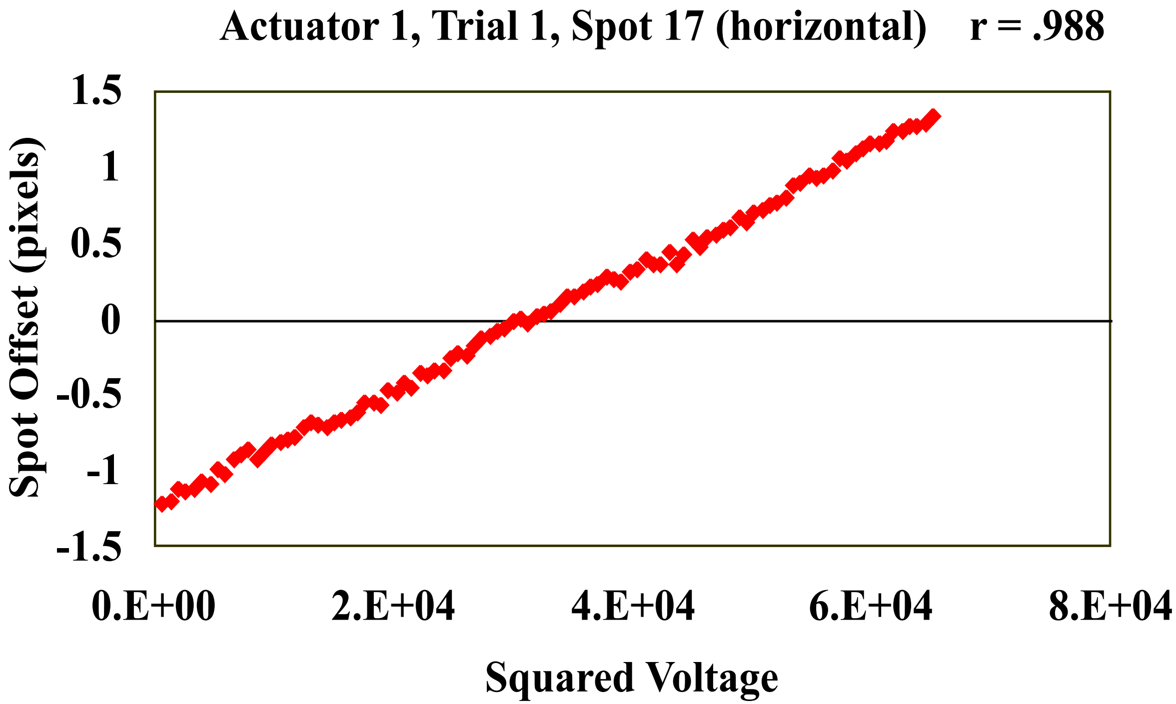
C = star positions

A = actuator positions

B = influence matrix describing influence of specific actuator position on star positions

Measuring the Influence Matrix

- measure centroid positions in subapertures for different settings of actuator *k*
- for actuator *k* and subaperture *n*, slope of best fit line is element (n, k) of influence matrix B



Determining the Control Vector

- Influence matrix B is known, C from wavefront sensor
- Find control vector A to correct for error in wavefront
- Matrix inversion of B?

$$A = B^{-1}C$$

- Overdetermined system
  - More centroid measurements than actuators
  - No exact solution A exists for given set of centroids
  - No exact B^{-1} exists (B is rectangular)
- Singular Value Decomposition: approximate B^{-1} in best possible way (minimizes wavefront error)

Typical AO Error Terms

- Fitting errors from insufficient approximation of the wavefront by the deformable mirror, mostly due to finite number of actuators

$$\sigma_{fit}^2 \approx 0.3\left(\frac{D}{r_0}\right)^{5/3}$$

- Temporal errors from time delay between measurement and correction, mostly due to exposure and readout time

$$\sigma_{temp}^2 \approx \left(\frac{t}{\tau_0}\right)^{5/3}$$

- Measurement errors from wavefront sensor

$$\sigma_{measure}^2 \sim S/N$$

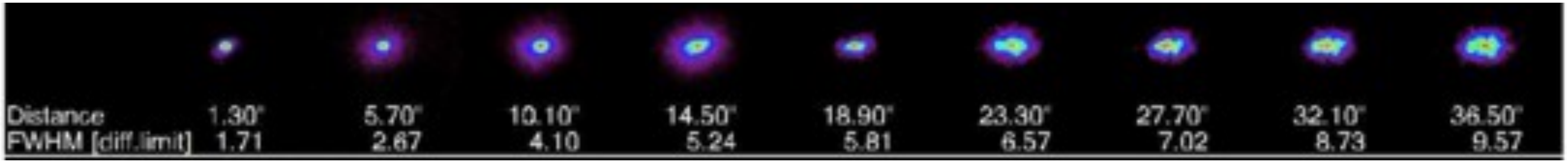
- Calibration errors from non-common aberrations between wavefront sensing optics and science optics

$$\sigma_{calibration}^2 \sim ???$$

- Angular anisoplanatism from sampling different lines of sight through the atmosphere, mostly limits field of view

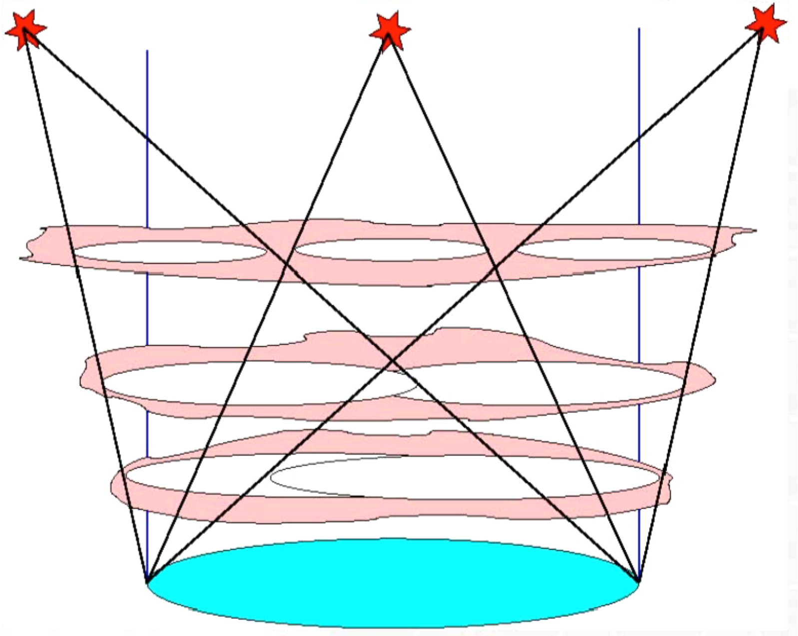
$$\sigma_{aniso}^2 \approx \left(\frac{\theta}{\theta_0}\right)^{5/3}$$

Angular Anisoplanatism

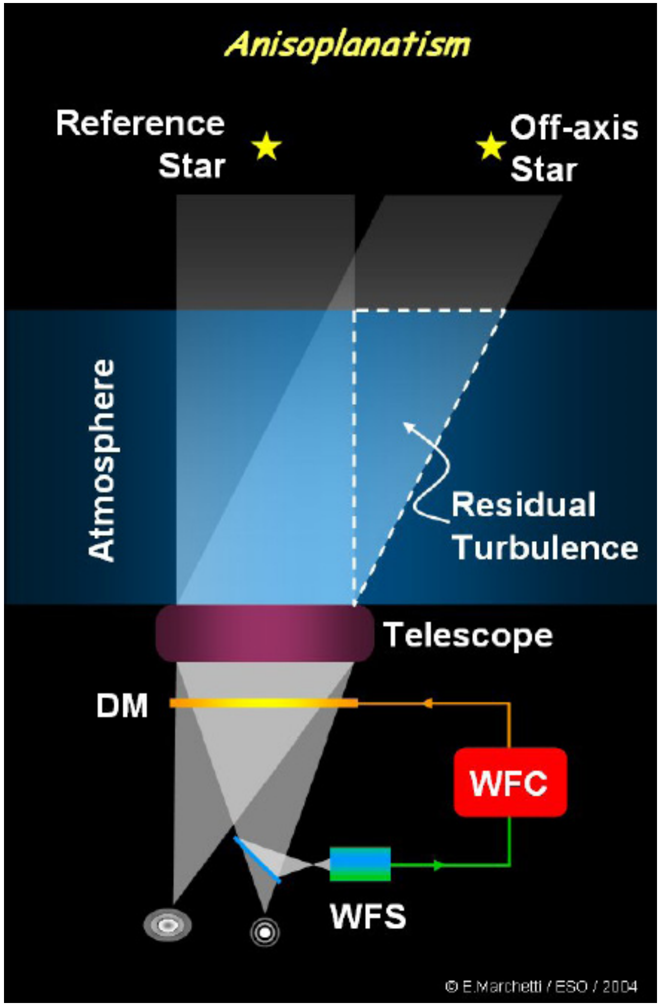


Angular anisoplanatism severely limits

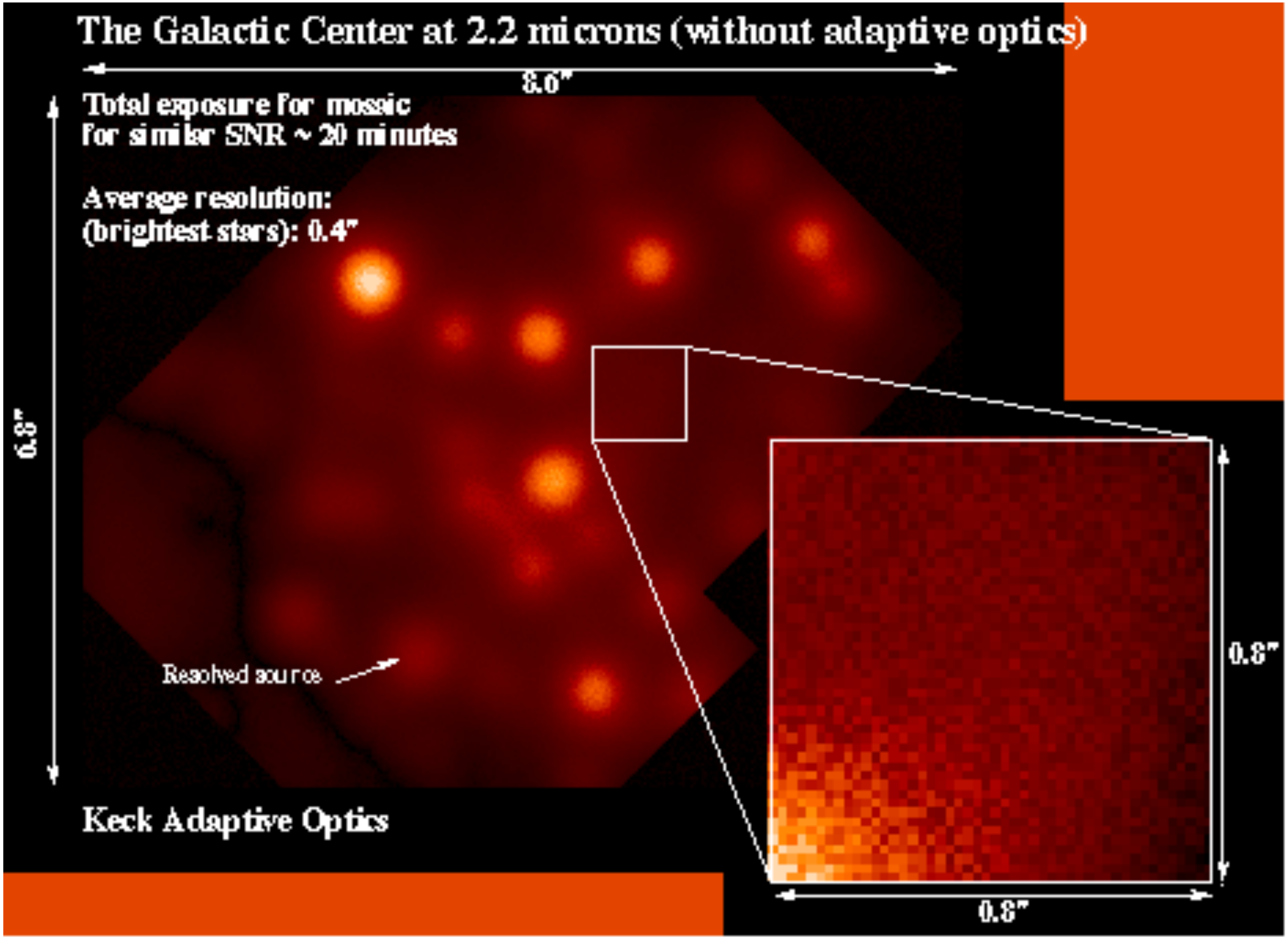
- wide-field imaging
- sky coverage (finding a guide star within the isoplanatic angle)



Multi-LGS allows to fight cone effect AND increase FOV



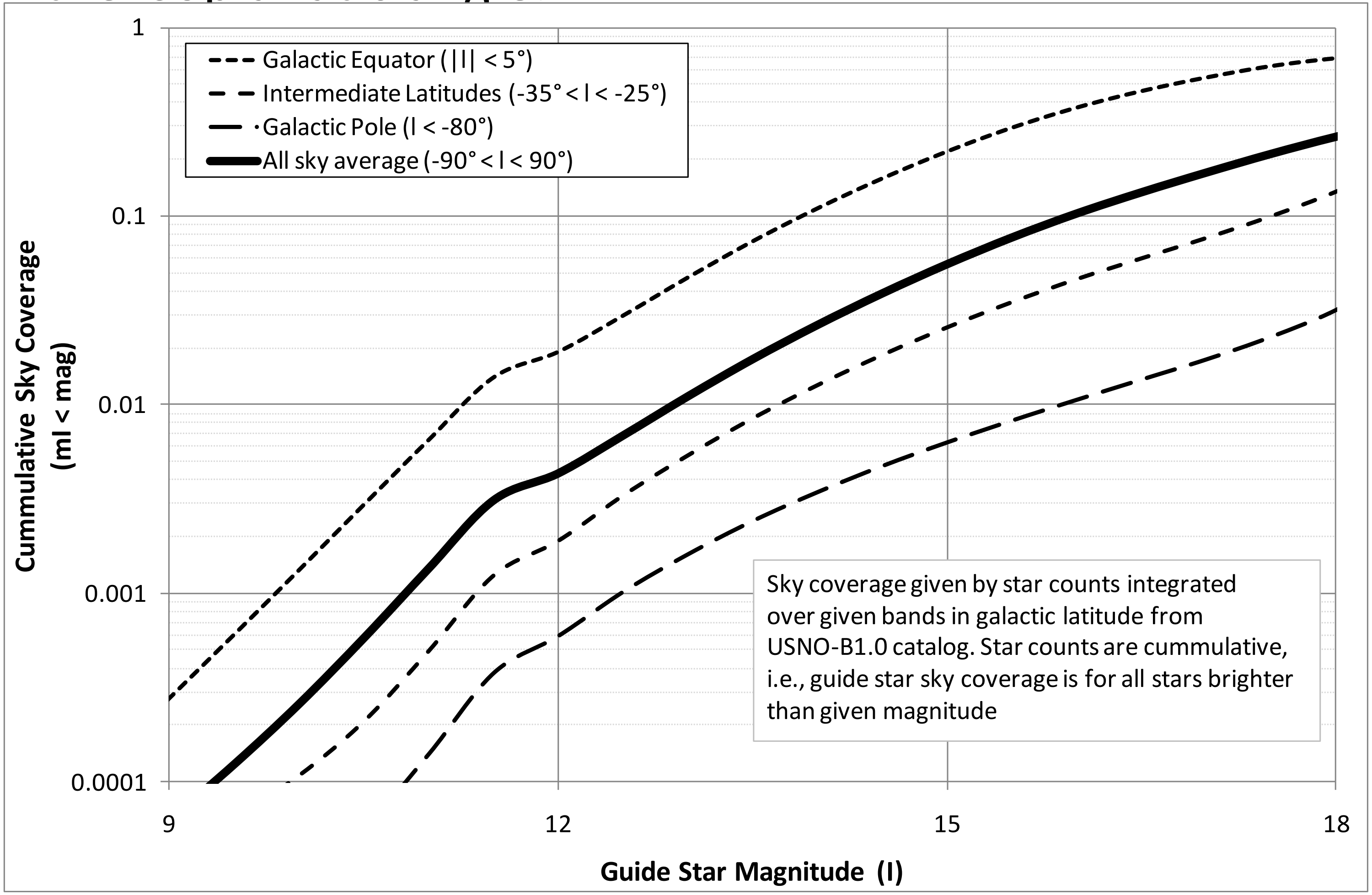
“Typical” Correction and Residuals



cfao.ucolick.org/pgallery/gc.php

Sky Coverage

To sense the wavefront one needs a bright reference/guide star within the isoplanatic angle.



Sky coverage given by star counts integrated over given bands in galactic latitude from USNO-B1.0 catalog. Star counts are cummulative, i.e., guide star sky coverage is for all stars brighter than given magnitude

Cumulative sky coverage, i.e., the chance of finding stars brighter than given magnitude, for a random target as a function of I-band magnitude using the USNO-B1.0 catalogue.

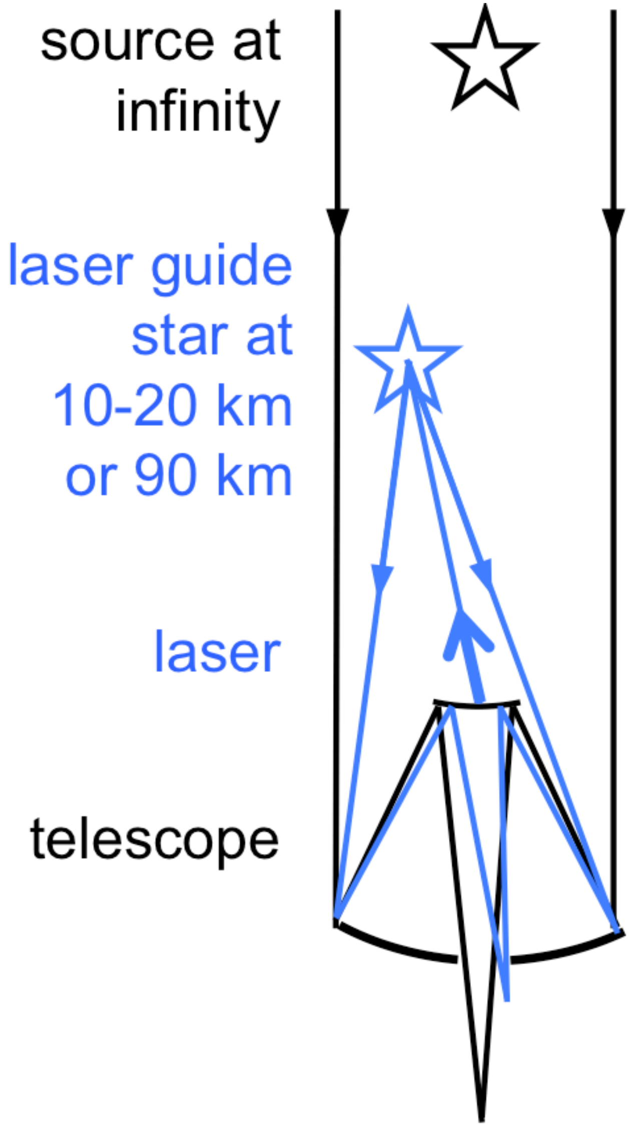
Laser Guide Stars

Solution to the sky coverage problem:

create your own guide star

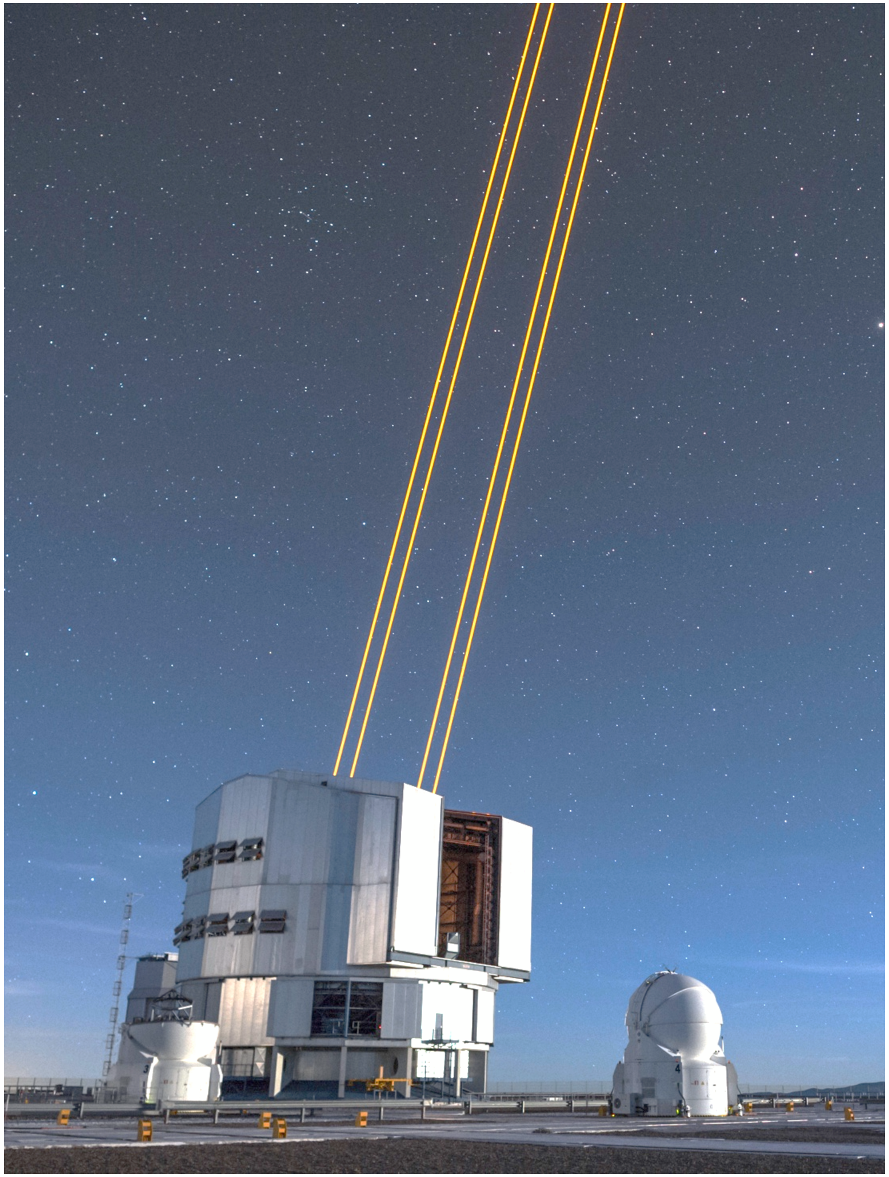
- Sodium LGS – excite atoms in "sodium layer" at altitude of ~ 95 km.
- Rayleigh beacon LGS – scattering from air molecules sends light back into telescope, h ~ 10 km

Since the beam travels twice (up and down) through the atmosphere, tip-tilt cannot be corrected → LGS-AO still needs a natural guide star, but this one can be much fainter (~18mag) as it is only needed for tip-tilt sensing



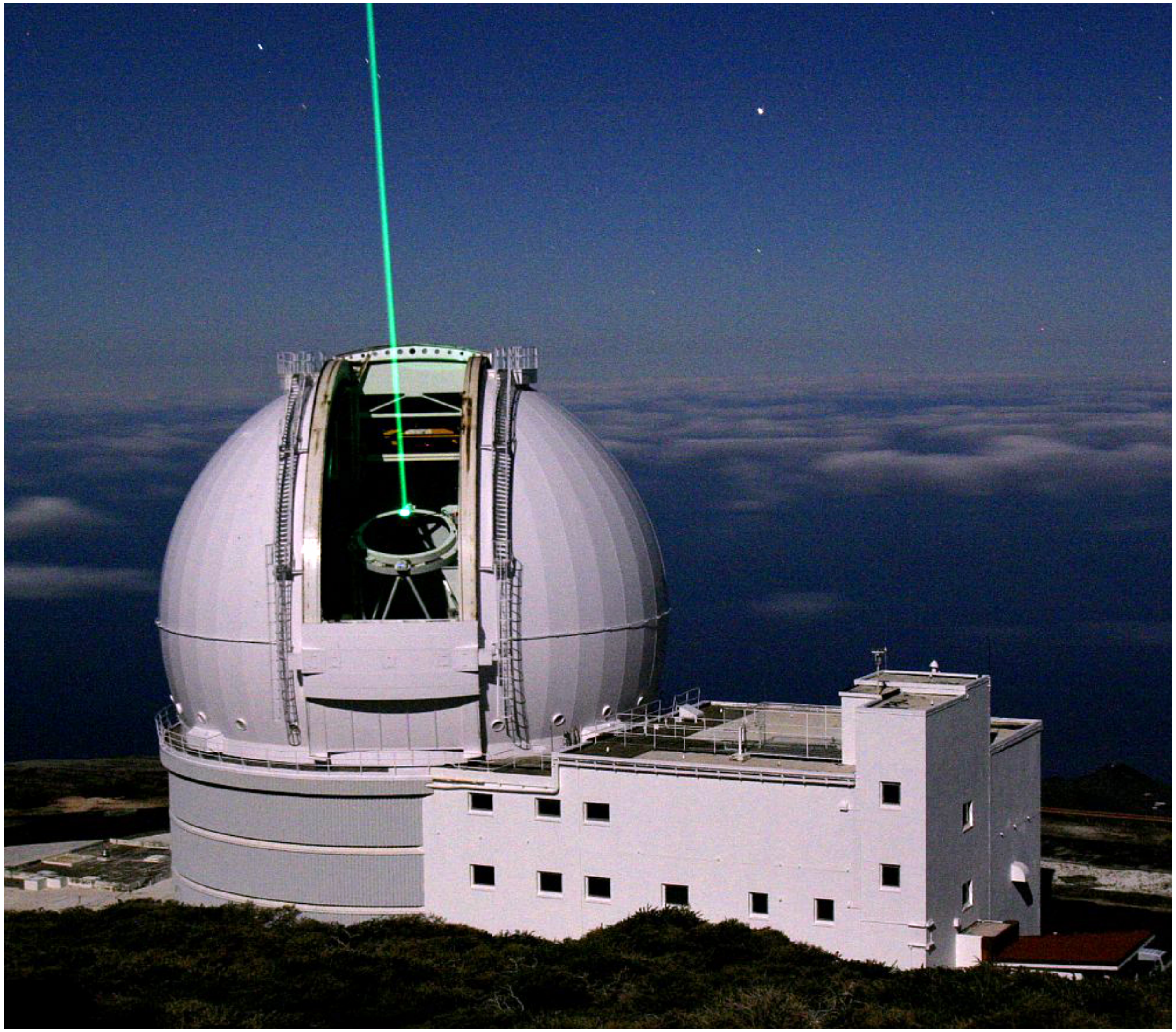
Sodium Beacons

- Layer of neutral sodium atoms in mesosphere (height ~ 95 km, thickness ~10km) from smallest meteorites
- Resonant scattering occurs when incident laser is tuned to D2 line of Na at 589 nm.



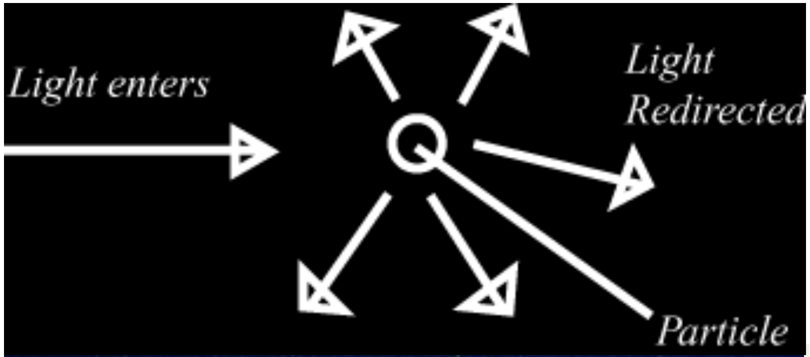
www.eso.org/public/images/eso1613n/

WHT Rayleigh Guide Star



Rayleigh Beacons

Due to interactions of the electromagnetic wave from the laser beam with molecules in the atmosphere.



Advantages:

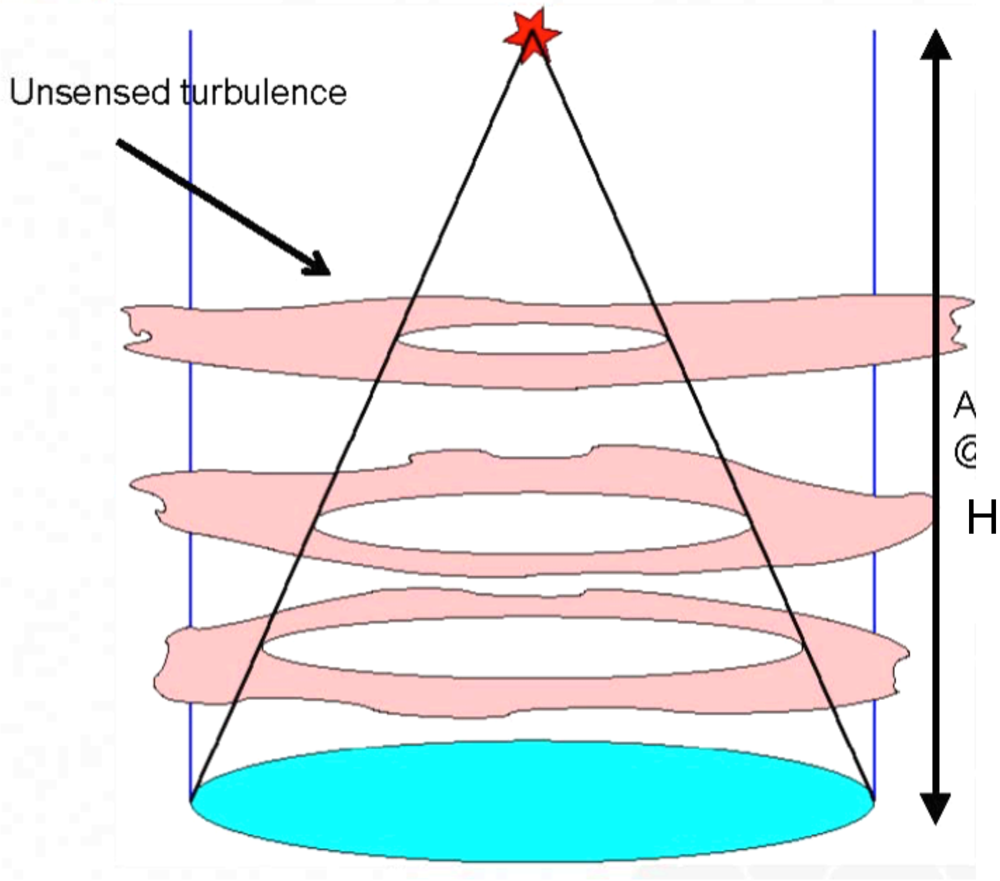
- cheaper and easier to build
- higher power
- independent of Na layer

Disadvantages:

- larger focus anisoplanatism
- laser pulses → timing

Focus Anisoplanatism

The LGS is at finite distance H above the telescope and does not sample all turbulence and not the same column of turbulent atmosphere ("cone effect"):



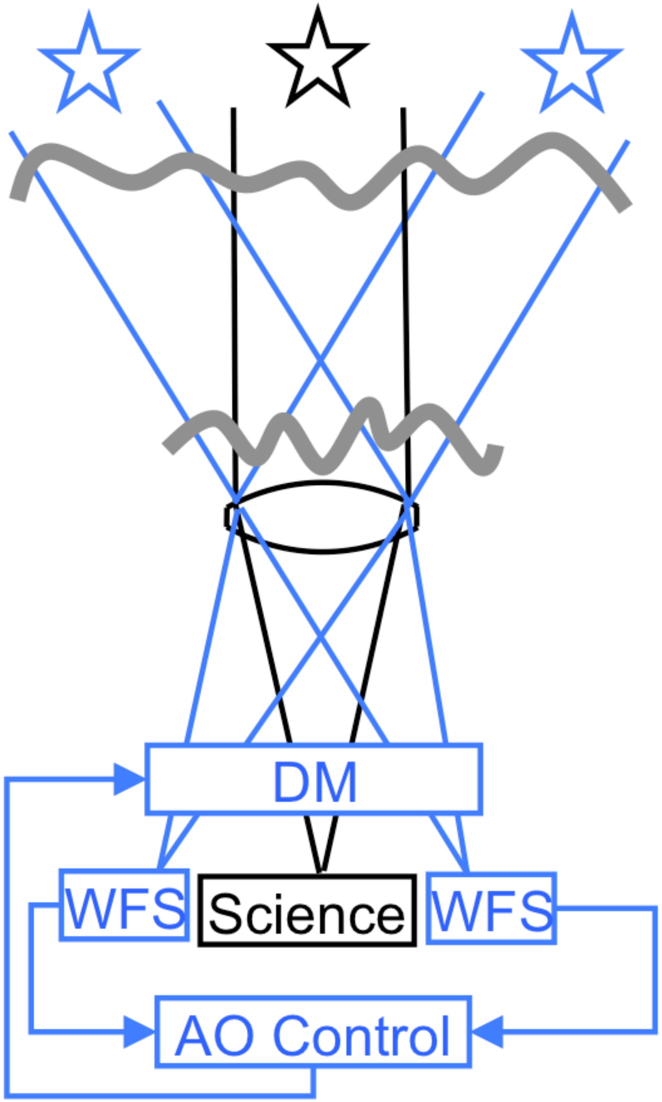
The contribution to the wavefront erro contribution from focus anisoplanatism is: $\sigma_{FA}^2 = \left(\frac{D}{d_0} \right)^{5/3}$

where $d_0 \sim \lambda^{6/5}$ depends only on wavelength and turbulence profile at the telescope site.

→ very large telescopes need multiple LGSs due to this cone effect.

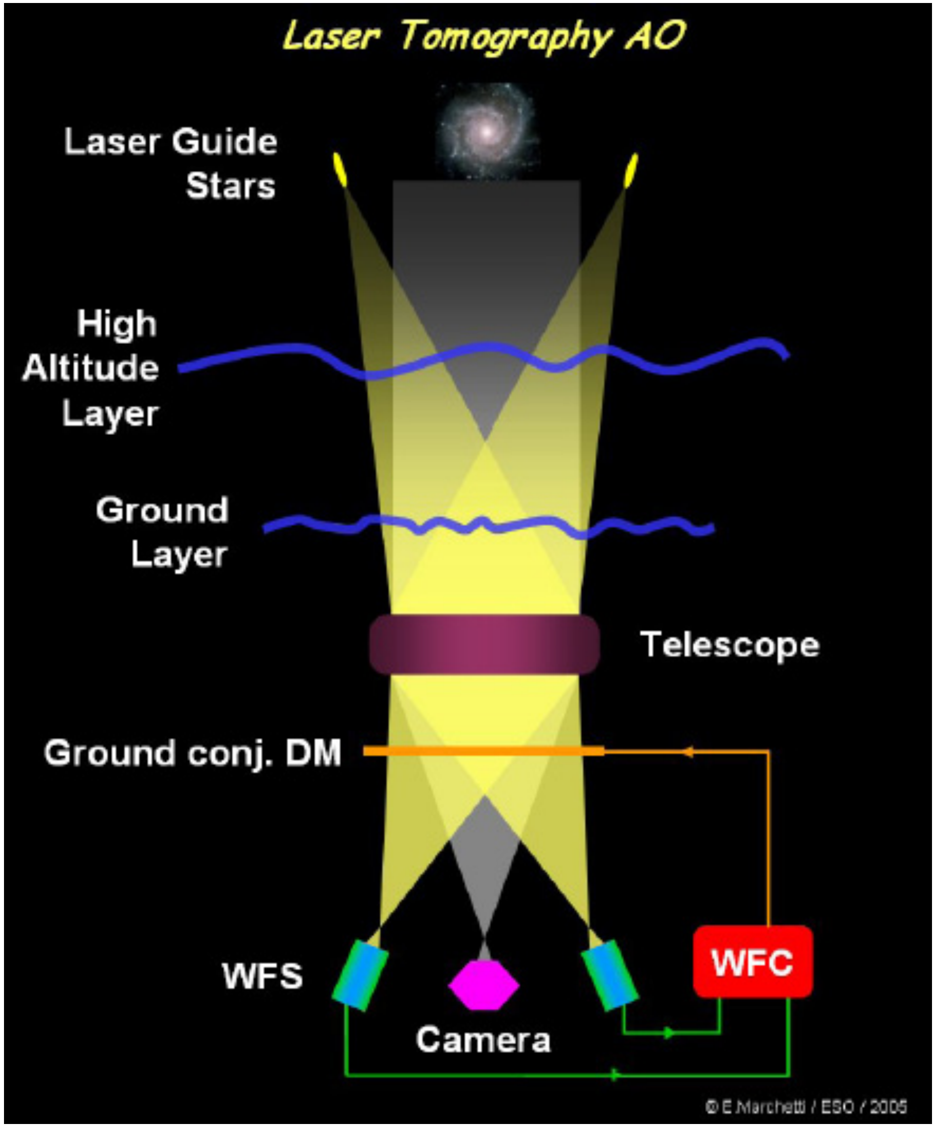
Ground Layer AO – GLAO

- Useful if ground layer (= ground + dome + mirror seeing) is the dominant component
- Uses several WFS and guide stars within a large FOV (several arcmin)
- WFS signals are averaged → control one DM
- Reduction of FWHM ~ factor of two (only!)
- GLAO is thus a "seeing enhancement" technique
- Advantage: wider fields and shorter wavelengths



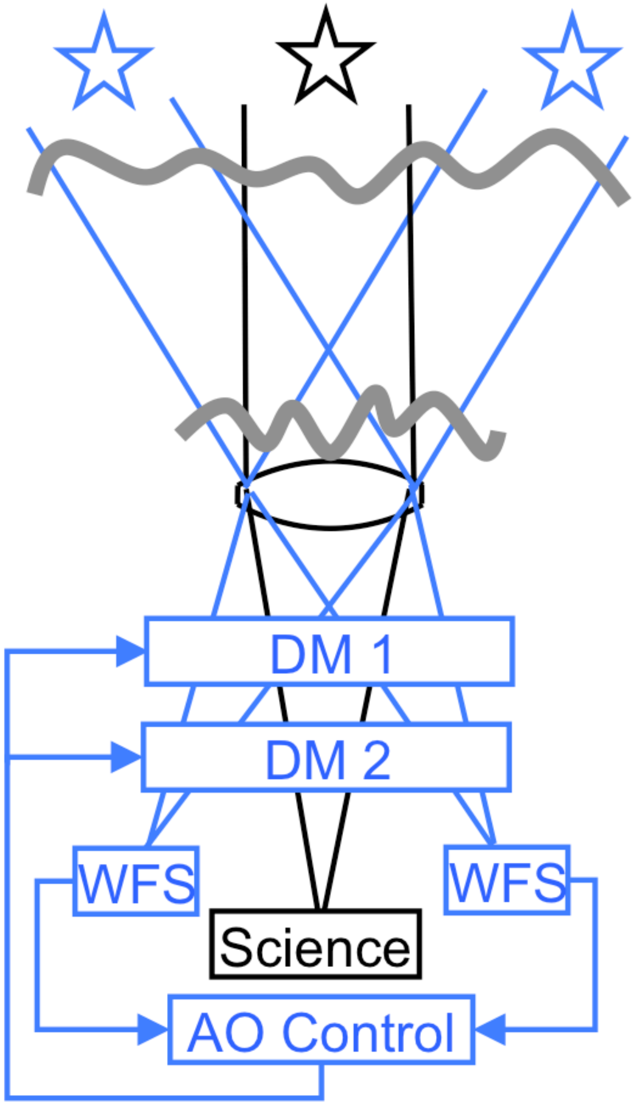
Laser Tomography AO – LTAO

- Uses multipe laser beacons
- each laser has its WFS
- combined information is used to optimize the correction by one DM on-axis.
- reduces the cone effect
- system performance similar to natural guide star AO but at much higher sky coverage.

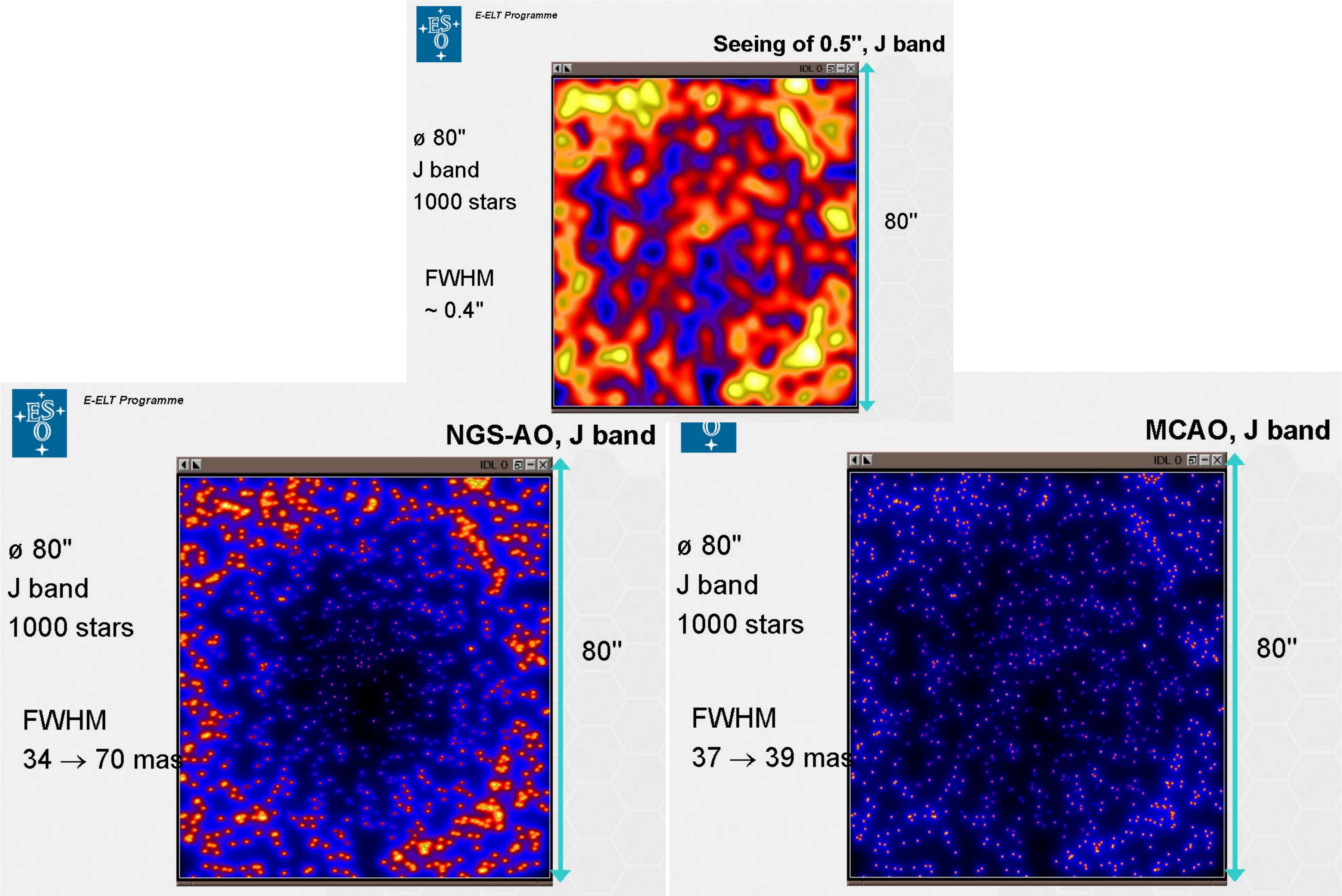


Multi-Conjugate AO – MCAO

- to overcome anisoplanatism, the basic limitation of single guide star AO
- MCAO uses multiple NGS or LGS
- MCAO controls several DMs
- each DM is conjugated to a different atmospheric layer at a different altitude
- at least one DM is conjugated to the ground layer
- best approach to larger corrected FOV



MCAO: Performance



E-ELT Programme

Seeing of 0.5", J band

ø 80"
J band
1000 stars

FWHM
~ 0.4"

80"

E-ELT Programme

NGS-AO, J band

ø 80"
J band
1000 stars

FWHM
34 → 70 mas

80"

MCAO, J band

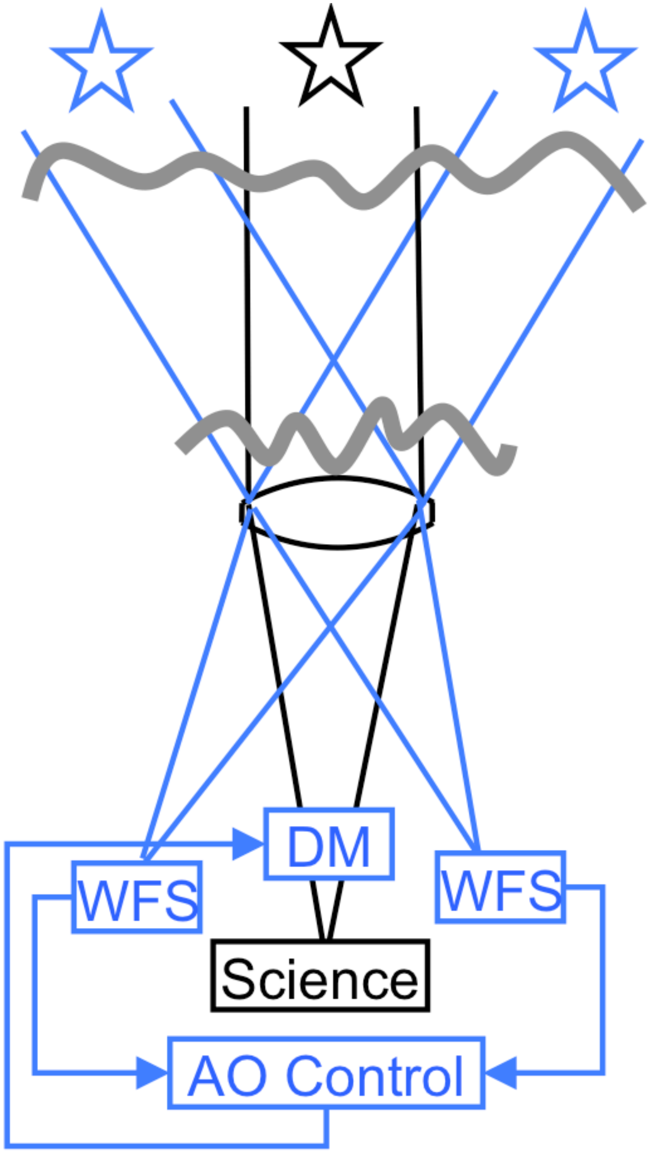
ø 80"
J band
1000 stars

FWHM
37 → 39 mas

80"

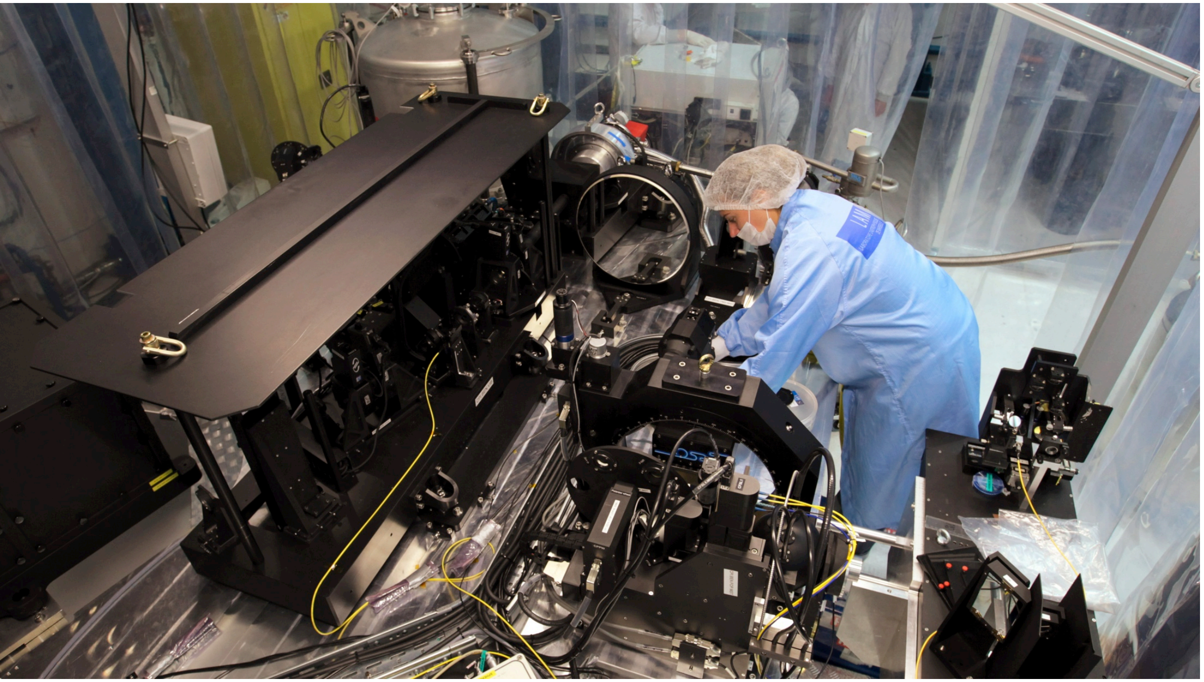
Multi-Object AO – MOAO

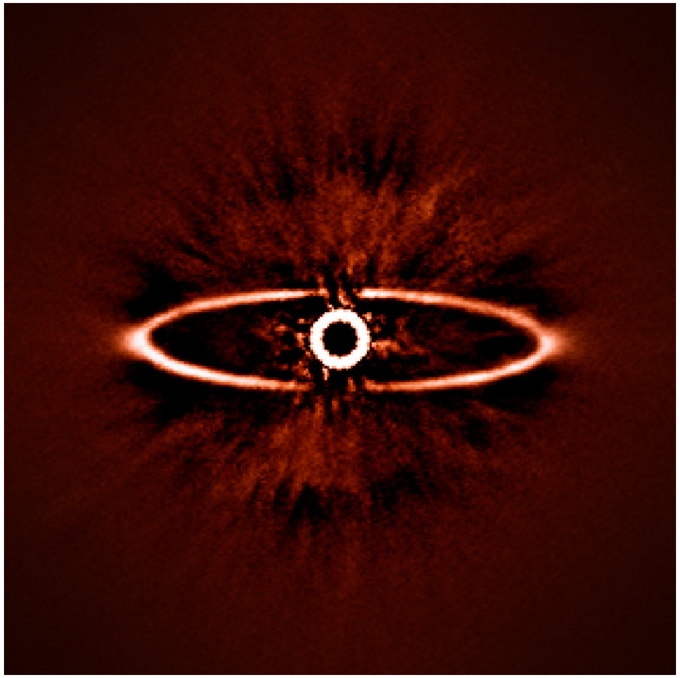
- MOAO provides correction not over the entire FOV of several arcmin but only in local areas within several arcmin → multi-object spectroscopy.
- needs (several) guide stars close to each science target.
- picks up the WFS light via small "arms“ inserted in the FOV.
- each science target has its DM
- systems work in open loop (!)



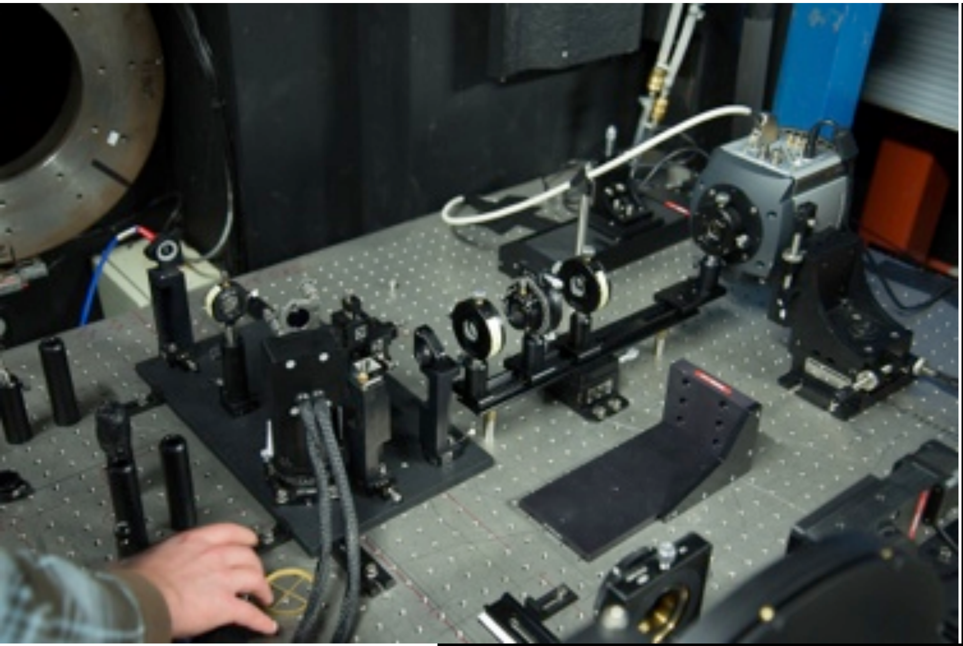
Extreme AO – XAO

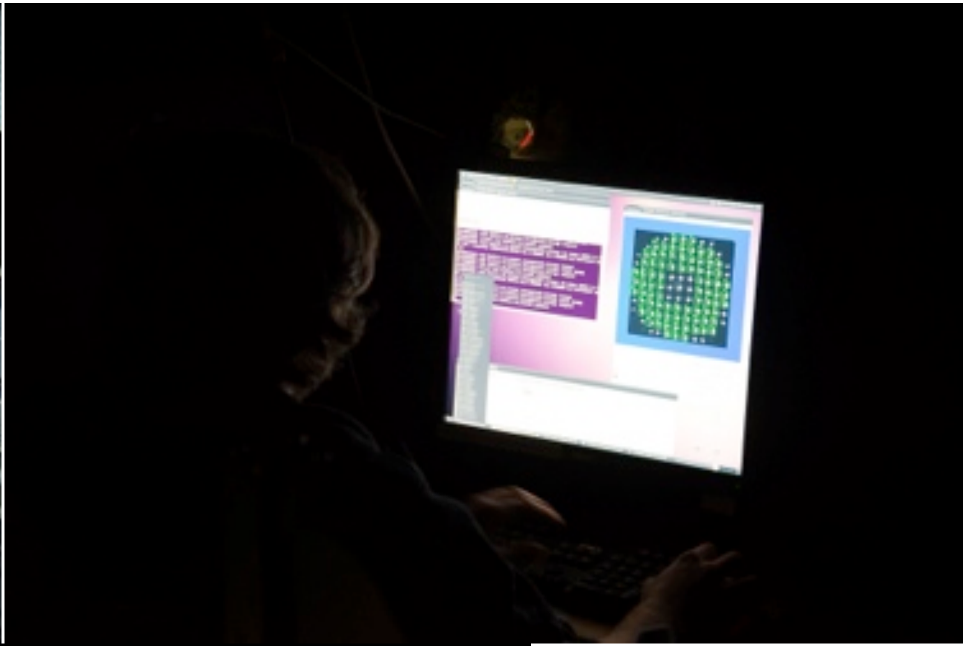
- XAO is similar to SCAO
- high Strehl on-axis and small corrected FOV
- however, Strehl values in excess of 90%
- requires many thousands of DM actuators
- requires minimal optical and alignment errors
- main application: search for exoplanets, SPHERE on VLT

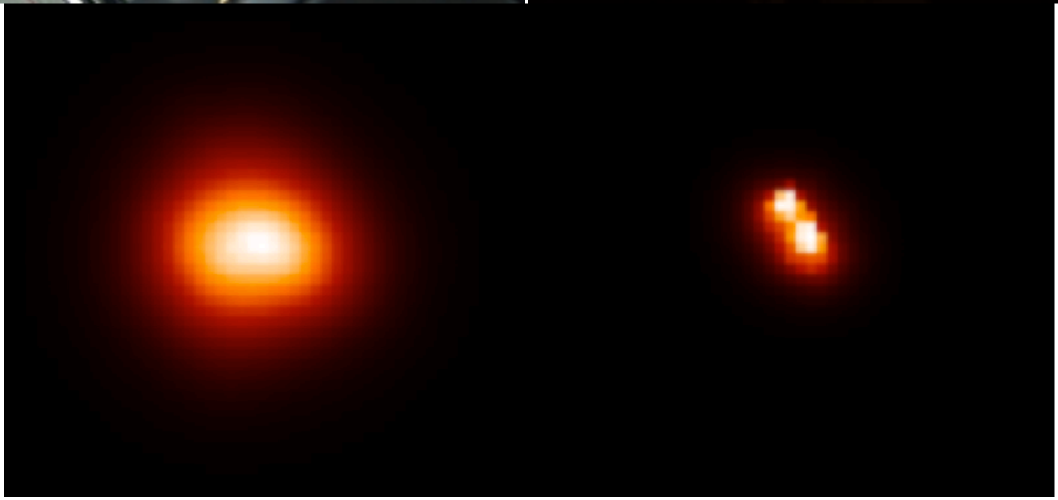




Student-Built ExPo Adaptive Optics







TODAY: Mercury Transit

- old observatory
- starts just before 14:00

Last Exercise Class: 12 May at 11:15

Presentations

- 26 May 2016, between 11:00 and 16:00
- 15 minutes per person
- location to be announced

MSc Astronomy & Instrumentation

- lectures:
  - Astronomical Telescopes and Instruments
  - Detection of Light
  - High Contrast Imaging
  - Astronomical Systems Design
  - Project Management
- option to take courses at TU Delft (not required)
- option for major research thesis in industry

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